

Bow-shock structure of Sgr-B molecular-cloud complex in the Galactic Centre inferred from 3D CO-line kinematics

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ABSTRACT

Three-dimensional (3D) bubble structure of the Sgr-B molecular-cloud complex is derived by a kinematical analysis of CO-line archival cube data of the Galactic Centre (GC) observed with the Nobeyama 45-m telescope. The line-of-sight depth is estimated by applying the face-on transformation method of radial velocity to the projected distance on the Galactic plane considering the Galactic rotation of the central molecular zone (CMZ). The 3D complex exhibits a conical-horn structure with the Sgr-B2 cloud located in the farthest end on the line of sight at radial velocity $v_{\text{lsr}} \sim 70 \text{ km s}^{-1}$, and the entire complex composes a lopsided bubble opening toward the Sun at $v_{\text{lsr}} \sim 50$ to 30 km s^{-1} . The line-of-sight extent of the complex is $\sim 100 \text{ pc}$ according to the large velocity extent for several tens of km s^{-1} from Sgr-B2 to the outskirts. The entire complex exhibits a flattened conical bubble with full sizes $\sim 40 \text{ pc} \times 20 \text{ pc} \times 100 \text{ pc}$ in the l , b and line-of-sight directions, respectively. Based on the 3D analysis, we propose a formation scenario of the giant molecular bubble structure due to a galactic bow shock, and suggest that the star formation in Sgr-B2 was enhanced by dual-side compression (DSC) of the B2 cloud by the Galactic shock wave from up-stream and expanding H II region from the down-stream side of the GC Arm I in Galactic rotation.

Key words: stars: formation – ISM: bubbles – ISM: clouds – H II region – ISM: molecules – Galaxy: centre.

1 INTRODUCTION

The Sgr-B molecular complex is the largest gaseous structure in the central molecular zone (CMZ) with total molecular-gas mass of a few $10^6 M_{\odot}$, where a remarkable ‘hole-and-ring’ structure around the radio sources Sgr-B1 and B2 has been found (Bally et al. 1987, 1988). The hole/ring structure has been interpreted in three ways:

- (i) Molecular-gas ring left after local perturbation by the large mass concentration of the Sgr-B2 cloud (Bally et al. 1988).
- (ii) Expanding molecular-gas bubble produced by the feedback of star-formation in Sgr-B1 and B2 (Sofue 1990).
- (iii) Cavity left after collision of the Sgr-B2 molecular cloud (Hasegawa et al. 1994; Sato et al. 2000; Enokiya & Fukui 2022).

The first model that (i) the hole was made by the gravitational perturbation in the complex has not been examined since the discovery. The second model, (ii) expanding-bubble and sequential star formation, was proposed as an alternative to (i) in our early paper (Sofue 1990), where we performed a detailed comparison of the spatial and kinematical distributions of molecular gas in the CO line and H II regions in radio continuum and recombination lines. We have shown that the molecular hole composes a shell structure at LSR (local standard of rest) velocity $v_{\text{LSR}} \sim 45 \text{ km s}^{-1}$ associated with Sgr-B1. We proposed a model that an expanding-bubble around Sgr-B1 hit the Sgr-B2 cloud at $v_{\text{LSR}} \sim 60\text{--}70 \text{ km s}^{-1}$ and triggered the recent and more active star formation in B2. The star-formation (SF) activity in Sgr-B2 further produced stellar winds and expanding

shock waves, which caused the larger-scale bubble in the extended molecular outskirts. More recent high-resolution CO-line mappings have shown that the complex is composed of a number of shell-like structures and high-density core clouds with large velocity extent (Oka et al. 1998; Tokuyama et al. 2019). Expanding shells and bubbles are more often recognized in the GC (Tsuboi, Ukita & Handa 1997; Tsuboi, Miyazaki & Okumura 2009; Tsuboi, Miyazaki & Uehara 2015).

The third idea in (iii) the cloud-collision paradigm has been extensively studied using molecular-line observational data, since the core-and-ring structure in the Sgr-B complex led to an idea that it may be a remnant of cloud collision (Hasegawa et al. 1994; Sato et al. 2000). Enokiya & Fukui (2022) recently performed multiwavelengths study and examined the model based on their criterion for the head-on collision of two clouds (Fukui et al. 2021). However, since the orbital elements of the two clouds, the necessary-and-sufficient condition for a head-on collision, have not been determined in the current papers reporting the ‘evidences’ for collisions, we may still have room to consider an alternative scenario.

In this paper, we explore the expanding-bubble model (ii) (Sofue 1990) in more detail, and show that the Sgr-B molecular complex can be reasonably explained by a galactic shock wave associated with an open cylindrical cavity caved from inside by an expanding H II region as the feedback of the star forming activity in Sgr-B1 and B2. New aspect in this work is that we take into account the Galactic rotation of the CMZ in the analysis of longitude–velocity diagrams (LVD), which has been ignored in the current studies. The spatial and velocity extents of the molecular complex, $\sim 50 \text{ pc} \times \sim 60 \text{ km s}^{-1}$, cannot be ignored compared to the radius $\sim 100 \text{ pc}$ and

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rotation speed $\sim 150 \text{ km s}^{-1}$ of the CMZ (Sofue 2013), because the rapid rotation with large differential velocity significantly affects the kinematics and evolution of a cloud with such a size. The spread longitude–velocity feature associated with Sgr-B complex manifests the Galactic rotation, leading to a largely extended distribution of the gas in the line-of-sight direction. Henshaw

As to the basic structure of the CMZ (Sofue 2022; Henshaw et al. 2023), we consider it to be a gaseous disc in Galactic rotation at $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$ in the deep gravitational well of the Galaxy (Sofue 2013). We consider that the molecular disc (CMZ) is superposed by spiral density waves, where the galactic shock-wave theory (Fujimoto 1968; Roberts 1969, 1972) applies to form dense molecular arms, GC Arm I and II (Sofue 1995). In the galactic shock waves, supersonic flow in galactic rotation encounters a stagnated gas in the arm rotating at the pattern speed, causing high density compression along the arm. In this paper, we assume that the GC Arm I and II compose the major structure of the CMZ, on which the SF regions Sgr-B1/B2 and Sgr C are located (Sofue 1990, 1995, 2022).

Here, Sgr-B1 and B2 are the thermal radio sources containing H II regions with their 110α -recombination-line coordinates at Sgr-B1: $(l, b, v_{\text{LSR}}, \Delta v) \sim (0^\circ.51, -0^\circ.05, +45.0 \text{ km s}^{-1}, 28.0)$, and B2: $(0^\circ.670, -0^\circ.036, +64.6 \text{ km s}^{-1}, 41.0 \text{ km s}^{-1})$, where Δv is the line width of the recombination line (Downes et al. 1980). The corresponding CO-line clouds are located at B1: $(l, b, v_{\text{LSR}}) \sim (0^\circ.52, -0^\circ.05, 30 - 40 \text{ km s}^{-1})$ and B2: $(0^\circ.66, -0^\circ.04, 60 - 70 \text{ km s}^{-1})$ as read in the maps in the following sections.

We consider a molecular-bow shock (MBS) in the spiral arm of CMZ, which is indeed observed in the Milky Way in the tangential direction of the 4-kpc molecular ring as W43 at G30 + 00 and at toward the Norma arm as G22 (Sofue et al. 2019; Sofue 2019). MBS is ubiquitous not only in the Galactic disc, but also in spiral galaxies and their central few hundred parsecs. In fact numerous MBS and giant cometary, H II regions (GCH) are found along the spiral arms of the barred spiral galaxy M83 (Sofue 2018) based on the optical image of M83 taken with the *Hubble Space Telescope* (McQuinn et al. 2017). Recent mid-infrared images of nearby spiral galaxies with the JWST (Lee et al. 2023; Barnes et al. 2023) have revealed numerous dust ‘voids’ in the face-on spiral galaxy M74, even in its central region, which are considered to be infrared views of MBS. It is stressed that the MBS + GCH structure is ubiquitous all over the galactic disc including the central region, which may also be the case in our Galactic Centre. In this paper, we focus on the Sgr-B molecular complex in the CMZ of our Galaxy in analogy to such MBS in the central molecular-gas discs of nearby galaxies.

In Section 2, we analyse the molecular gas distribution and kinematics in the ^{13}CO -line data cube. Section 3 is devoted to construction of three-dimensional (3D) molecular maps of the Sgr-B complex by applying the face-on transformation (FOT) method to the rotating Galactic Centre disc. In Section 4, we propose three possible mechanisms to explain the observed bubble property of Sgr-B2 complex. In Section 5, we discuss the result in comparison with the current models, and summarize the paper in Section 6. We assume the distance to the Galactic Centre to be $R_0 = 8.2 \text{ kpc}$ (GRAVITY Collaboration et al. 2019)

2 MOLECULAR-CLOUD PROPERTY OF SGR-B COMPLEX

2.1 Data

We used the archival data cube of the ^{13}CO ($J = 1 - 0$) line emission at 110.27 GHz taken from the CMZ survey of $3^\circ \times 0.7$ region using

the Nobeyama 45-m telescope (Tokuyama et al. 2019). The telescope had a FWHM (full width of half maximum) angular resolution of 15 arcsec , velocity resolution of 1 km s^{-1} , rms (root mean squared) noise temperature 1 K. The data are re-gridded to a cube with $(7.5 \text{ arcsec} \times 7.5 \text{ arcsec} \times 2 \text{ km s}^{-1})$ sampling with an effective resolution of 16.7 arcsec . In the analysis, we use the brightness temperature, T_B , which is assumed to be equal to the radiation temperature, T_R , used in the archival data, because the gaseous structures discussed in this paper are extended and well resolved by the beam. We also used partially the archival cube data of the HCN ($J = 4 - 3$) line emission at 354.5 GHz from the GC survey with ASTE (Atacama Sub-mm Telescope Experiments) 10-m telescope (Tanaka et al. 2018). The telescope had angular resolution of 22 arcsec , and the data are resampled into $(8.5 \text{ arcsec} \times 8.5 \text{ arcsec} \times 2 \text{ km s}^{-1})$ grids with the effective spatial resolution of $24''$.

2.2 Integrated intensity (moment 0) map: the mass and conversion factor

Moment maps of the ^{12}CO and ^{13}CO line intensities are shown in Fig. 1. The major and minor axial diameters of the ^{12}CO moment 0 map are measured to be $D_x \sim 78.7 \text{ pc}$ and $D_y \sim 32.1 \text{ pc}$, respectively, yielding the size radius of $r = \sqrt{D_x D_y} / 2 \sim 25.1 \text{ pc}$. Adopting a CO-to- H_2 conversion factor of $X_{^{12}\text{CO}} = 0.5 \times 10^{20} \text{ H}_2 \text{ cm}^{-2} [\text{K km s}^{-1}]^{-1}$ and mean molecular weight $\mu = 1.38$ (Arimoto, Sofue & Tsujimoto 1996), we have $M_{\text{mol}} = \pi r^2 I_{^{12}\text{CO}} X_{^{12}\text{CO}} (2m_{\text{H}} \mu) \sim 1.7 \times 10^6 M_\odot$.

The velocity width is measured to be $w_1 \sim 12 \text{ km s}^{-1}$ on the moment 2 map, which is nearly uniform over the cloud. The velocity gradient observed in the moment 1 map adds another velocity width due to internal motion by rotation, expansion and/or contraction of $w_2 = \sqrt{(v_{\text{max}} - v_{\text{min}})^2} \sim 10 \text{ km s}^{-1}$, where $v_{\text{max}} \sim 52$ and $v_{\text{min}} \sim 32 \text{ km s}^{-1}$ as read on the moment 1 map. Thus, we obtain an effective velocity dispersion of $v_\sigma = \sqrt{w_1^2 + w_2^2} \sim 15.6 \text{ km s}^{-1}$. The dynamical (virial) mass is then estimated by $M_{\text{vir}} = r v_\sigma^2 / G \sim 1.4 \times 10^6 M_\odot$. The corresponding kinetic energy is $E_k \sim 1/2 M_{\text{vir}} v_\sigma^2 \sim 3.4 \times 10^{51} \text{ erg}$.

Therefore, the molecular mass is about 1.24 times the dynamical mass, and the cloud is bound by the gravity. However, if we assume that the cloud is virialized, or the gravitational force is balanced with the dynamical pressure by the velocity dispersion, so that $M_{\text{vir}} = M_{\text{mol}}$, we can determine the conversion factor to be $X_{^{12}\text{CO}} \sim 0.4 \times 10^{20} \text{ H}_2 \text{ cm}^{-2} [\text{K km s}^{-1}]^{-1}$. This value is consistent with that predicted by the gradient of $X_{^{12}\text{CO}}$ with the galacto-centric distance as found for general values in the disc (Arimoto et al. 1996).

2.3 Velocity field and dispersion (moment 1 and 2) maps showing conical outflow toward the observer

Moment 1 map, or the velocity field, is shown in Fig. 1, which reveals negative velocity gradients from the centre toward the edge of the cloud. The radial velocity decreases from $v_{\text{LSR}} \sim 70 \text{ km s}^{-1}$ in the central red clump toward the eastern edge at $\sim 40 \text{ km s}^{-1}$ (green outskirts) at velocity gradient of $dv/dx \sim 0.4 \text{ km s}^{-1} \text{ pc}^{-1}$, and at $dv/dx \sim 0.8 \text{ km s}^{-1} \text{ pc}^{-1}$ toward the western edge at 30 km s^{-1} (blue outskirts).

Moment 2 map shows that the dispersion velocity attains minimum of $\sigma_v \sim 8 \text{ km s}^{-1}$ in the central molecular clump that is surrounded by an extended cloud with larger velocity dispersion of $\sigma_v \sim 12 \text{ km s}^{-1}$. These velocity structures can be interpreted as due to an accelerating conical gas flow in the direction of the observer (Sun), blue-shifting from the central red clump toward the outer envelope. This structure can be more clearly demonstrated by a composite colour map shown

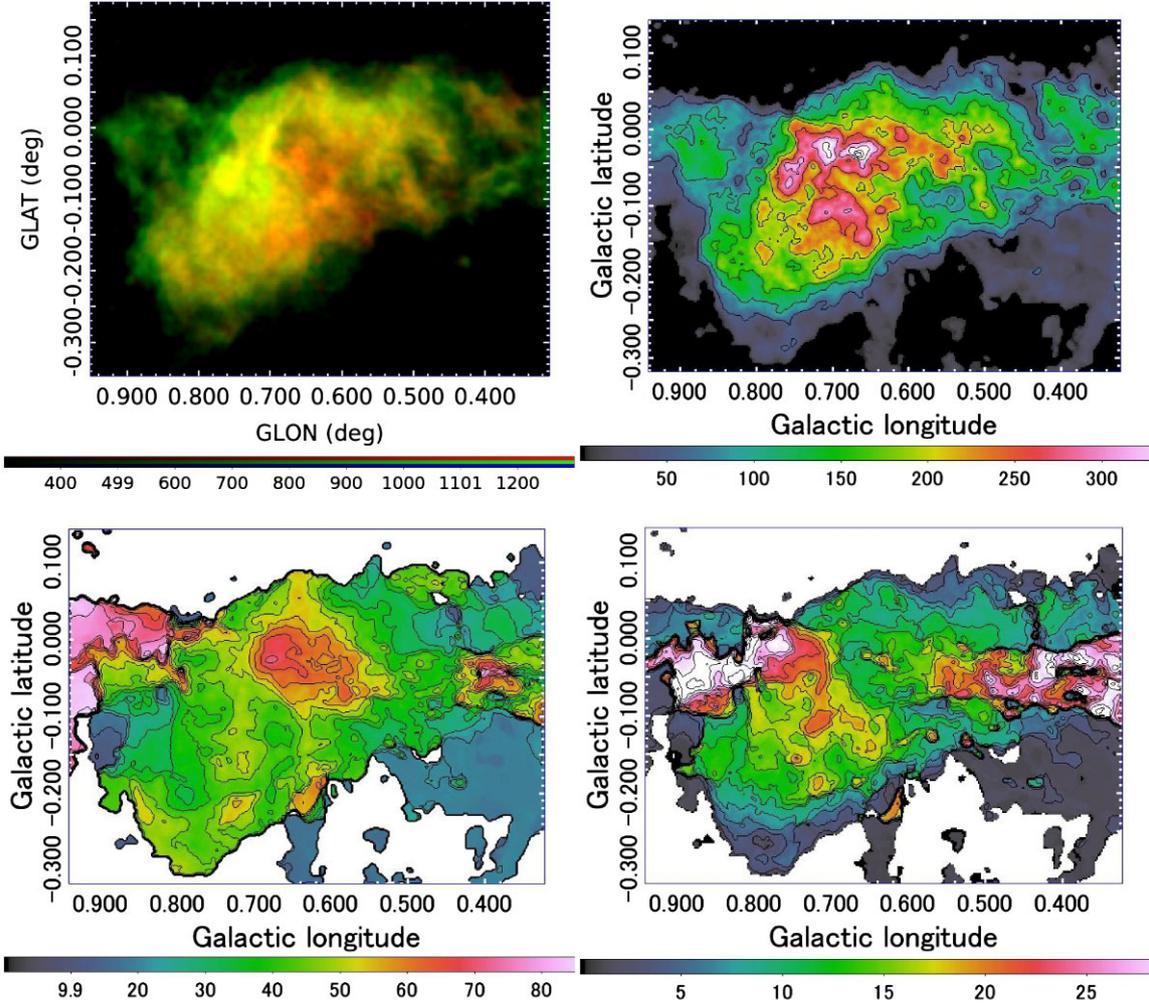


Figure 1. [Top left] Overlay of integrated intensity from $v_{\text{LSR}} = 30$ to 70 km s^{-1} of the ^{12}CO line emission in green and ^{13}CO in red of the Sgr-B molecular cloud. [Top right] Moment 0, [bottom left] moment 1 and [bottom right] moment 2 in ^{13}CO line integrated from $v_{\text{LSR}} = 10$ to 90 km s^{-1} . The horizontal broad feature is the contamination of Arm II at $v_{\text{LSR}} \sim 80 \text{ km s}^{-1}$. Note the difference from top-left panel, where the integration is from 0 to 70 km s^{-1} in order to avoid contamination of Arm II.

in Fig. 2, where the three different-velocity clumps are overlaid with $v_{\text{LSR}} = 71 \text{ km s}^{-1}$ in red colour, 51 km s^{-1} in green, and 31 km s^{-1} in blue.

2.4 Channel maps: cavity, shell, and radio continuum sources

Fig. 3 shows channel maps of Sgr-B molecular components overlaid with contours of the radio continuum emission at 1.28 GHz from the MeerKAT observations (Heywood et al. 2022; Yusef-Zadeh et al. 2022). The figure demonstrates the spatial correlation of the radio sources Sgr-B1 and B2 with the molecular cavity in the extended components of Sgr-B molecular cloud.

Fig. 2 shows overlay of the brightness at 75 km s^{-1} in red, 40 km s^{-1} in green and 35 km s^{-1} in blue. The red clump near the cloud centre composes a compact and dense central clump associated with Sgr-B2, and is embedded in the green and blue components widely spread around the centre clump. As a whole, the Sgr-B complex composes a conical horn structure opening from the central compact red clump in the GC-Arm I toward the green and blue-shifted outskirts.

The radio continuum source Sgr-B1 associated with a supernova remnant is located in touch with the western wall of the cone making

the green-component of the extended cloud in Fig. 2. It is stressed that the western wall is observed as an arc-shaped high-density molecular arc showing up at $\sim 30\text{--}50 \text{ km s}^{-1}$ as shown in the upper panel of Fig. 3.

The Sgr-B2 radio continuum source is full of point-like sources, mostly compact H II regions, surrounded by extended H II diffuse emission, and is located in the highest-density molecular clump recognized as the red component, as shown in Fig. 3.

2.5 LVDs, BVDs, and kinematics

The velocity structure is quantitatively represented by position-velocity diagrams (PVD). Fig. 4 shows longitude-velocity diagrams (LVD) at various latitudes and latitude-velocity diagrams (BVD) at various longitudes of Sgr-B complex. Fig. 5 shows LVD in wider longitudinal range, and enlargement near Sgr-B complex. The LVDs reveal that the main structure of GC Arm (GCA) I runs as a straight LV ridge running from $(l, v) \sim 0^\circ.3, 20 \text{ km s}^{-1}$ to $(0^\circ.8, +100 \text{ km s}^{-1})$ as indicated by the dashed line in Fig. 5 (A).

The outskirts composed of the green (intermediate, $\sim 50 \text{ km s}^{-1}$) and blue-shifted ($\sim 30 \text{ km s}^{-1}$) components of the complex are

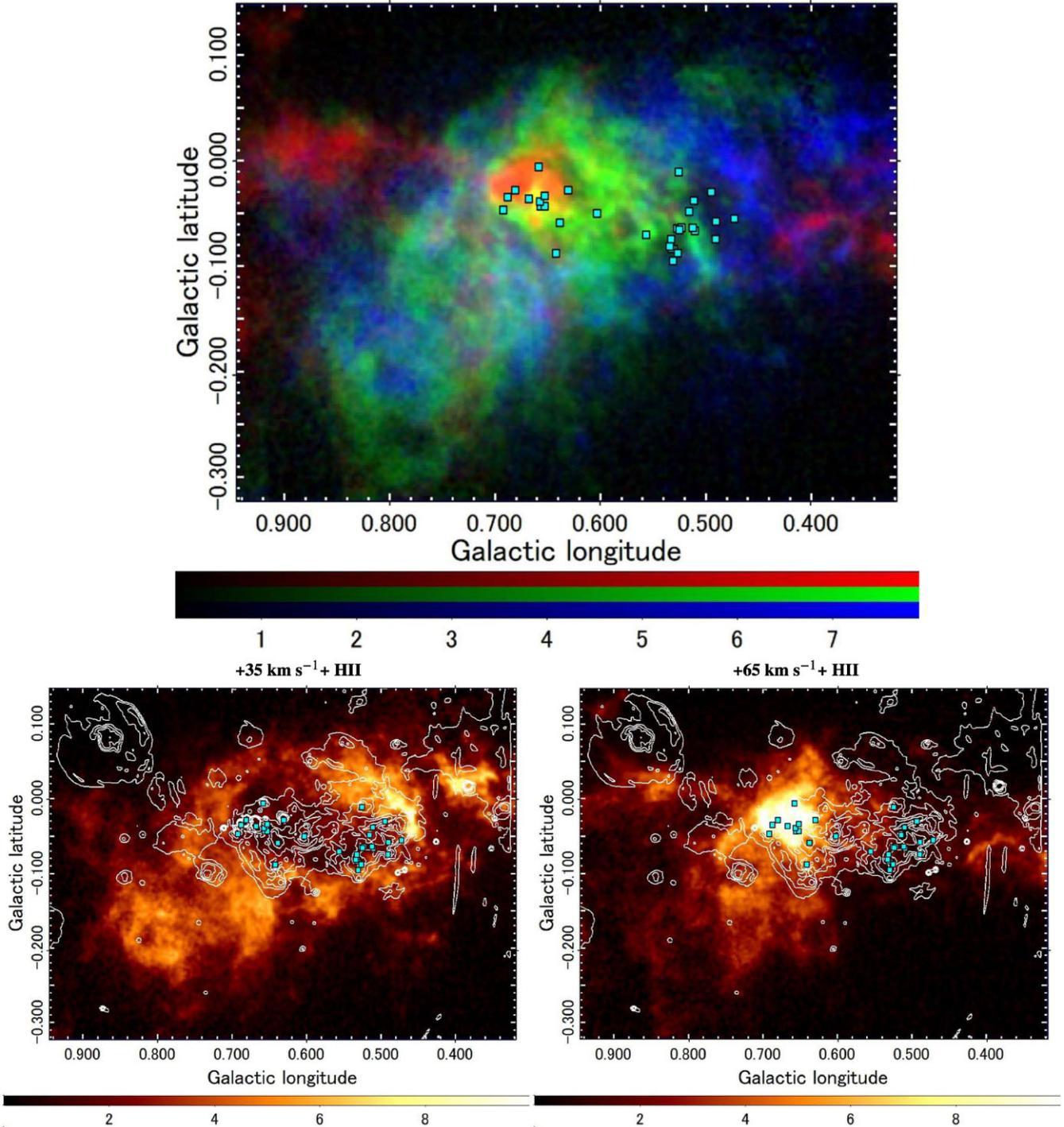
$^{13}\text{CO } T_B$ at $v_{\text{LSR}} = +75$ (red), 49 (green), & 35 km s^{-1} (blue) + HII regions

Figure 2. Overlay map of the brightness temperature T_B (scaled by the bars in K) at $v_{\text{LSR}} = 35$ (in blue), 49 (green), and 75 km s^{-1} (color only online). Rectangles are positions of H II regions detected in the 110α recombination line (Mehring et al. 1992, 1993). The cluster at $l \sim 0^\circ.67$ toward the red molecular clump is the Sgr-B2 H II region, and that at $l \sim 0^\circ.51$ is Sgr-B1. [Bottom] Same, but at $v_{\text{LSR}} = +35$ (left: as blue in the top panel) and 65 km s^{-1} (right: as red). The contours are 1.3 GHz continuum surface brightness by MEERKAT (Heywood et al. 2022; Yusef-Zadeh et al. 2022). Rectangles are the same as the top panel.

recognized as an open bow-shaped ridge in the LVD as traced by the dashed arc in panel (B) of Fig. 5. Thus, the molecular gas composing Sgr-B1, B2 and the outskirts of the complex are smoothly connected on the LVD. This fact indicates that they compose a continuous bow structure, or there is no signature of discrete

structures suggesting two different clouds that are colliding at high speed.

Panel (C) shows the latitude-velocity diagram (BVD) across Sgr-B2. This BVD shows a clearer ‘V’ shape, which is typical for a stagnated bow flow seen along the flow axis (Wilkin 1996). In the case

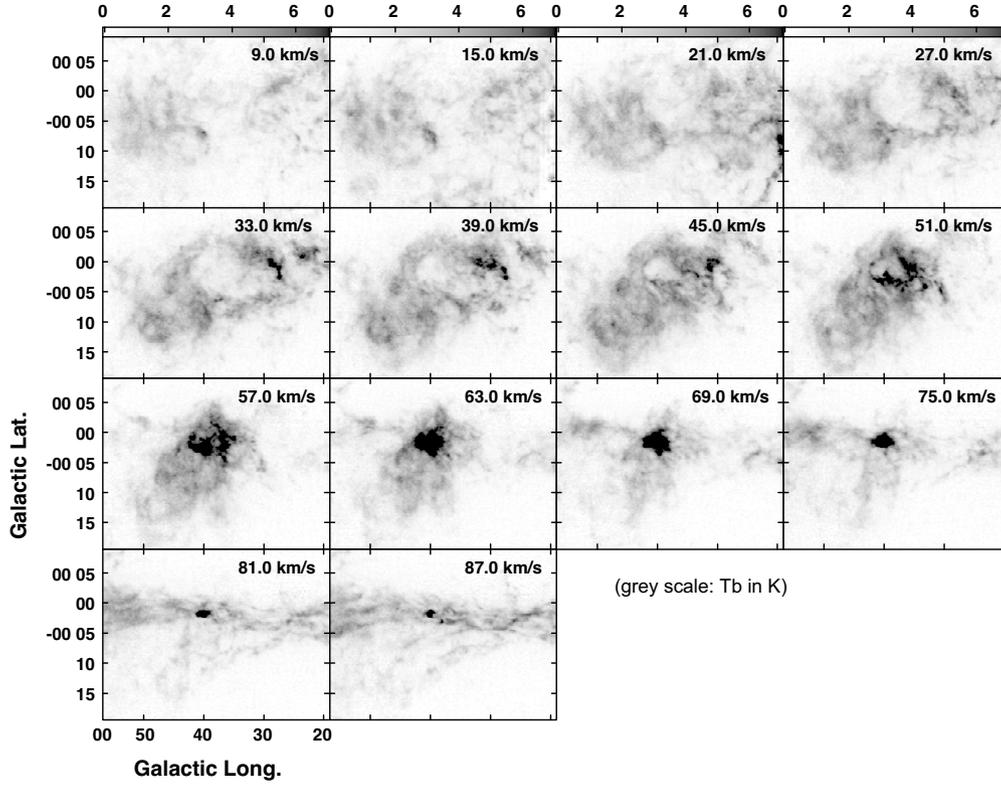


Figure 3. Channel maps of $^{13}\text{CO } T_B$ in K in order to demonstrate that the cavity surrounded by a shell (bubble) and the Sgr-B2 cloud are continuous structure in the velocity direction.

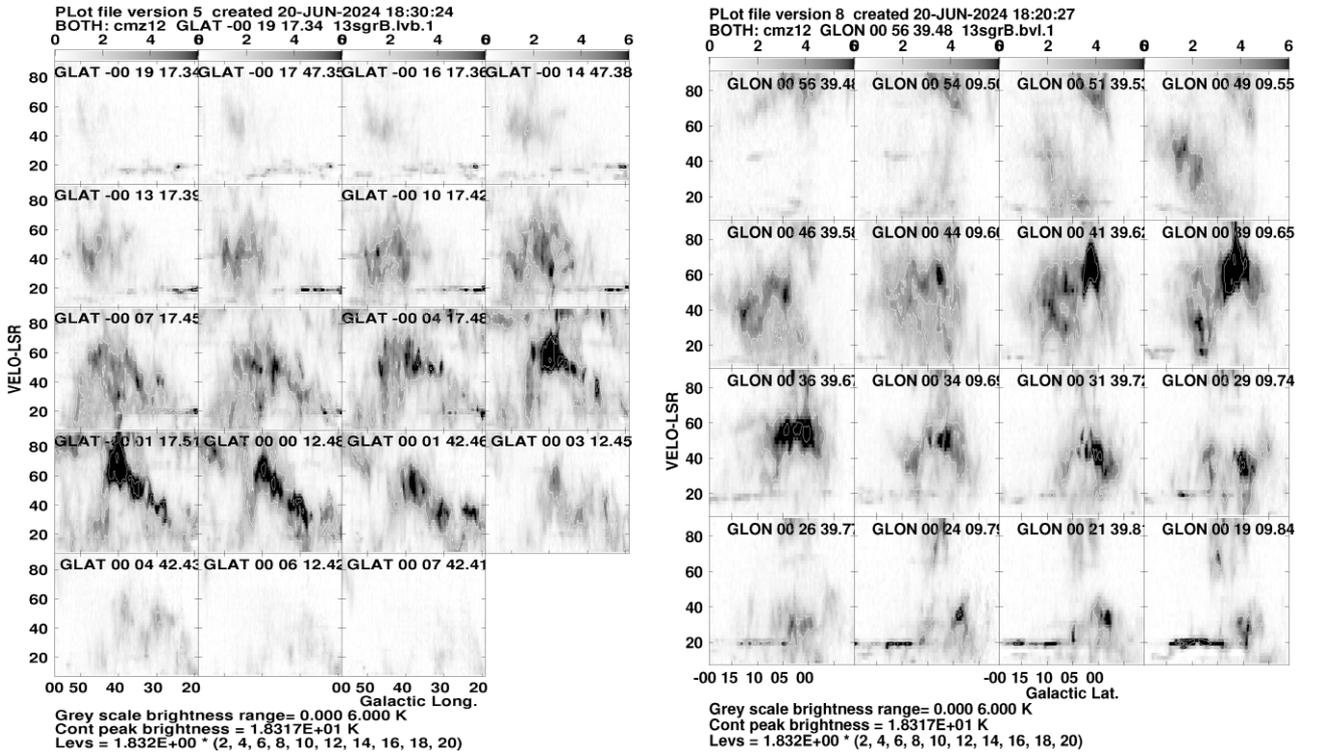


Figure 4. [left] LVD (Longitude–Velocity diagrams) at various b . [Right] BVD at various l .

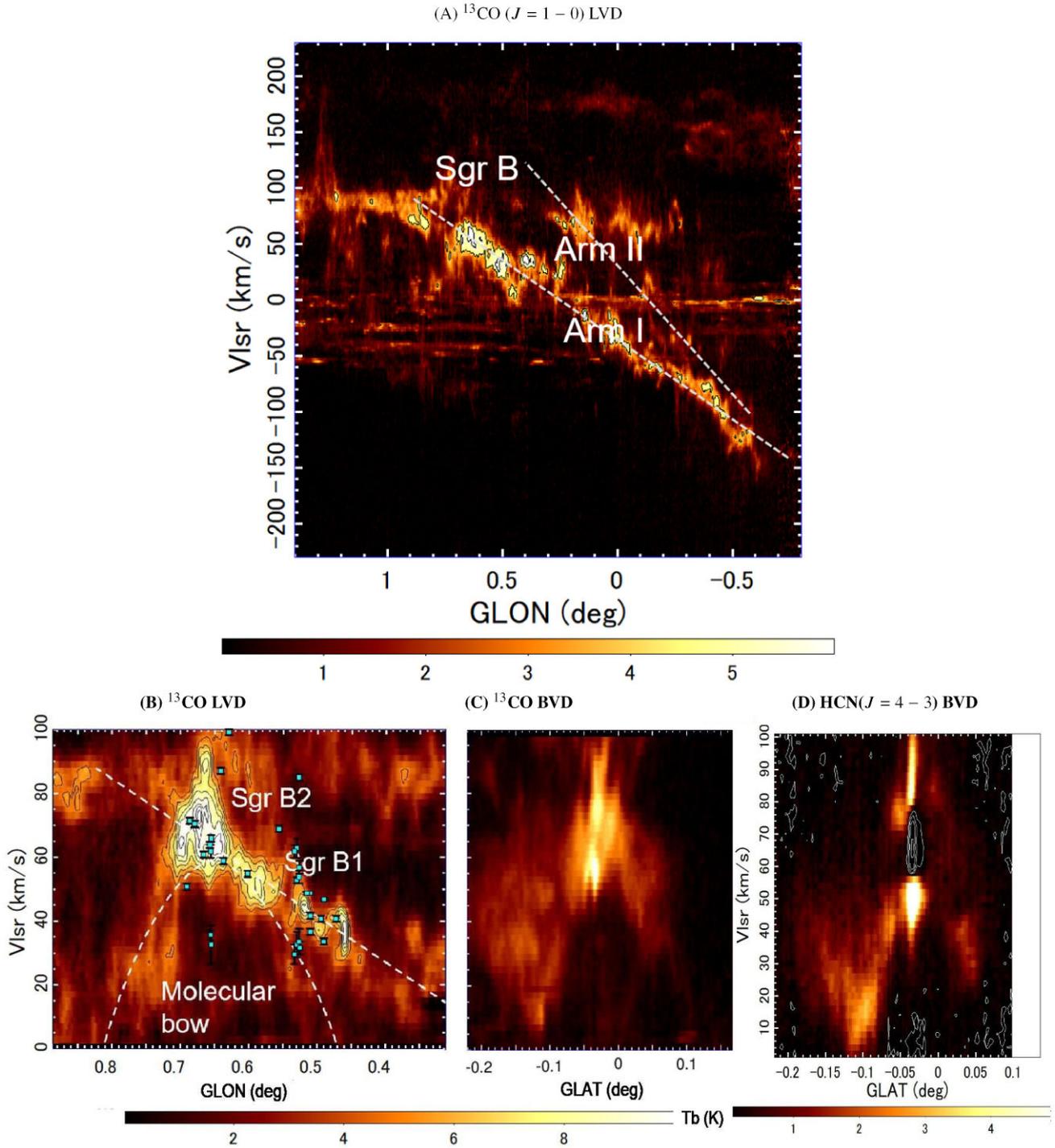


Figure 5. (A) ^{13}CO LVD along the entire ridge of GC Arm I at $b = +0^\circ.02$. Contours are every 1 K of T_{B} of the ^{13}CO line emission. The Sgr-B cloud composes the major structure of the Arm, from which the $30\text{--}50\text{ km s}^{-1}$ cloud extends to compose a bow shape open toward decreasing velocity. (B) LVD across Sgr-B2 at $(l, b, v_{\text{LSR}}) \sim (0^\circ.66, -0^\circ.04, 65\text{ km s}^{-1})$ and Sgr-B1 ($0^\circ.52, -0^\circ.05, 35\text{ km s}^{-1}$). Blue rectangles with error bars are H110 α recombination-line sources of H II regions (Mehringer et al. 1992, 1993). Sgr-B1 and B2 and associated molecular clouds are located exactly on the LV ridge of the GC Arm I (straight dashed line), indicating that they are rotating with the Arm around the GC. The curved dashed line indicates the extended outskirts at $20\text{--}50\text{ km s}^{-1}$, composing a bow structure (conical cylinder) in the LVD. (C) BVD perpendicular to the Galactic plane across Sgr-B2, exhibiting a V (Δ) shape typical for a stagnated bow flow (Wilkin 1996) toward the down-stream (low velocity) side of the Galactic rotation. Full width of the bow is $(\Delta X \times \Delta Z \times \Delta v_{\text{LSR}}) \sim (40\text{ pc} \times 20\text{ pc} \times 60\text{ km s}^{-1})$, significantly narrower than the l -directional width in the left panel. (D) BVD in the HCN ($J = 4 - 3$) line, the high-density tracer of molecular density $\gtrsim 10^7\text{ H}_2\text{ cm}^{-3}$ (Tanaka et al. 2018). The gas is highly concentrated along the V-shaped bow shock with the peak at $v_{\text{LSR}} \sim 66\text{ km s}^{-1}$ on the stagnation point. The line is saturated at Sgr-B2 and exhibits absorption against the continuum of Sgr-B2 as represented by white contours from -0.1 to -2.1 K every 0.5 K . Abscissa scales in the bottom three panels are approximately equal, showing that the vertical bow width is narrower than horizontal width.

of Sgr-B2, the flow is toward the down-stream side of the Galactic shock wave at the GC Arm I. Panel (D) shows the same, but in the HCN ($J = 4-3$) line, as extracted from the survey of high-density tracers in the sub-mm wave lines in the GC with ASTE (Atacama Sub-mm Telescope Experiment) (Tanaka et al. 2018). The V-shaped bow features both in the LVD and BVD across Sgr-B2 indicate a flattened corn structure in the (l, b, v_{LSR}) space. The full widths of the bow are measured to be $(\Delta X \times \Delta Z \times \Delta v_{\text{LSR}}) \sim (40 \text{ pc} \times 20 \text{ pc} \times 60 \text{ km s}^{-1})$. The V-shaped feature will be discussed later in more detail in a bow-shock model.

2.6 Tight alignment of Sgr-B1 and B2 clouds and H II regions along the GC Arm I

H II regions detected in the 110α recombination-line are marked by blue rectangles on the T_{B} maps in Fig. 2 and on a LVD in Fig. 5 with error bars (Mehring et al. 1992, 1993).

Recombination-line sources associated with Sgr-B1 composes a cluster around $(l, b, v_{\text{LSR}}) \sim 0^\circ.51, -0^\circ.04, 35 \text{ km s}^{-1}$ near the western edge of the molecular bubble at $v_{\text{LSR}} \sim 35 \text{ km s}^{-1}$. Those associated with Sgr-B2 compose a cluster around $(0^\circ.67, -0^\circ.02, +65 \text{ km s}^{-1})$ tightly overlapped with the massive cloud (B2 cloud, or the red clump).

Sgr-B1 and B2 molecular clouds and H II regions are located exactly along the main LVD ridge of the GC Arm I. The exact alignment of the B1 and B2 clouds and H II regions in space and velocity with the Arm indicates that the molecular clouds nesting Sgr-B1 and B2 are moving with the Arm without significant displacement from the Galactic rotation.

3 FACE-ON TRANSFORMATION (FOT) AND 3D STRUCTURE

3.1 FOT

In this section, we derive the three-dimensional (3D) distribution of molecular gas, assuming that the local velocities such as due to the expanding-bubble motion on the order $\sim 10 \text{ km s}^{-1}$ are sufficiently smaller than that of the galactic rotation on the order of $\sim 150 \text{ km s}^{-1}$. Because the rotation curve in the GC region at galacto-centric distances $R \sim 120 \text{ pc}$ is nearly flat at $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$ (Sofue 2013), we here assume that the gas in the CMZ is circularly rotating at a constant velocity of 150 km s^{-1} .

The 3D position (X, Y, Z) of a pixel with CO brightness temperature $T(v)$ at (l, b) can be transformed from the observed position and velocity (l, b, v_{LSR}) as

$$X = R_0 \sin l \quad (1)$$

$$Z = R_0 \tan b, \quad (2)$$

and

$$Y \simeq X \left(\frac{v}{V_{\text{rot}}} \right)^{-1} \sqrt{1 - \left(\frac{v}{V_{\text{rot}}} \right)^2}. \quad (3)$$

3.2 3D structure

Figs 6 shows the result of FOT for the brightness clipping above $T_{\text{B}} = 2 \text{ K}$. In the figure the small dots are plotted so that the density of dots is proportional to the volume density of the molecular gas. The top-left panel of Fig. 6 is a projection of the gas-density distribution on the sky or on the (X, Z) plane, representing the integrated column

density, which is therefore identical to the moment 0 map of the CO line emission. The top-right panel is the projection on the (Y, Z) plane, or a view from the right side (from eastern side). The bottom-left panel shows the projection on the (X, Y) plane, or on the Galactic plane, so that it expresses the face-on view of the complex. The magenta dots represents a ring of radius 120 pc in the galactic plane, approximately tracing the GC Arm I. The bottom-right panel is the same, but projected on the (X, Y) plane (galactic plane) in a wider area including the whole giant molecular cloud (GMC). Fig. 7 shows the same, but with the same linear scaling in the $X, Y,$ and Z directions with the intensity presented by grey scale in arbitrary unit as indicated by the bar.

Fig. 8 shows oblique projections of the plots viewed from a standpoint in the north-eastern and nearer side to the Sun. The two (left and right) panels show the same, but projections from different standpoints rotated by 10 deg from each other, so that the two panels can be used to obtain a stereo-graphic view of the gas distribution by viewing the panels by the two eyes.

In order to see how the representative parts of the Sgr-B complex according to the radial velocity corresponding to the resulting 3D distributions, we show in the left panel of Fig. 9 the same plots, but giving three different colours red, green, and blue (RGB) to the points having far, intermediate and near Y values. Namely, each coloured region represents a slice perpendicular to the line of sight. Similarly in the right panel of Fig. 9, we show the same plots, but RGB colours represent points with high, intermediate, and low LSR velocities of the ^{13}CO -line emission. So, this diagram represents equal-velocity slices, each representing the 3D distribution of each of the RGB components in Fig. 2.

Figs 6 to 9, thus, show that the molecular gas is densest in the Sgr-B2 clump at $v_{\text{LSR}} \sim 70 \text{ km s}^{-1}$, and the extended outskirts compose conical cylinder at velocities of ~ 50 to 30 km s^{-1} , which is open toward positive Y -direction (toward the Sun) and is blue shifted with respect to Sgr-B2 densest clump. It is stressed that the line-of-sight (Y) extent of the Sgr-B complex, from Sgr-B2 cloud to the blue-shifted outskirts, is as long as $\Delta Y \sim 100 \text{ pc}$. Namely, the complex composes a conical bubble of 3D (X, Y, Z) sizes of

$$\Delta X \times \Delta Y \times \Delta Z \sim 40 \text{ pc} \times 100 \text{ pc} \times 20 \text{ pc}. \quad (4)$$

The line-of-sight size is longest, which is due to the large velocity width of $\Delta v_{\text{LSR}} \sim 60 \text{ km s}^{-1}$.

3.3 Uncertainty due to expanding motion and velocity offset

In the present FOT, we have ignored the expanding motion of the molecular gas. If we correct for such local motion, the Y -directional extent of the 3D distribution is reduced, so that the bubble shape becomes more round and compact. However, a possible velocity correction of $\sim 10 \text{ km s}^{-1}$ by the expansion of H II region is much smaller than the observed velocity extent of the entire Sgr-B complex, $\sim 60 \text{ km s}^{-1}$, due to the Galactic rotation. So, the correction by H II region expansion can reduce only ~ 20 percent at most in the Y -direction. This means that the large velocity difference from $v_{\text{LSR}} \sim 20$ to $\sim 70 \text{ km s}^{-1}$ in the Sgr-B complex is due to the Galactic rotation, as simulated by a spiral structure formed in a disc in differential rotation (Kruijssen et al. 2019). We comment that such stationary bubble/conical-horn structure may be similar to the dust voids observed in the spiral galaxies by JWST (Lee et al. 2023; Barnes et al. 2023).

Another aspect to be mentioned is the uncertainty due to v_{LSR} displacement from the rotation centre on the LVD. As equation 3 indicates, the Y directional distance depends on the radial velocity

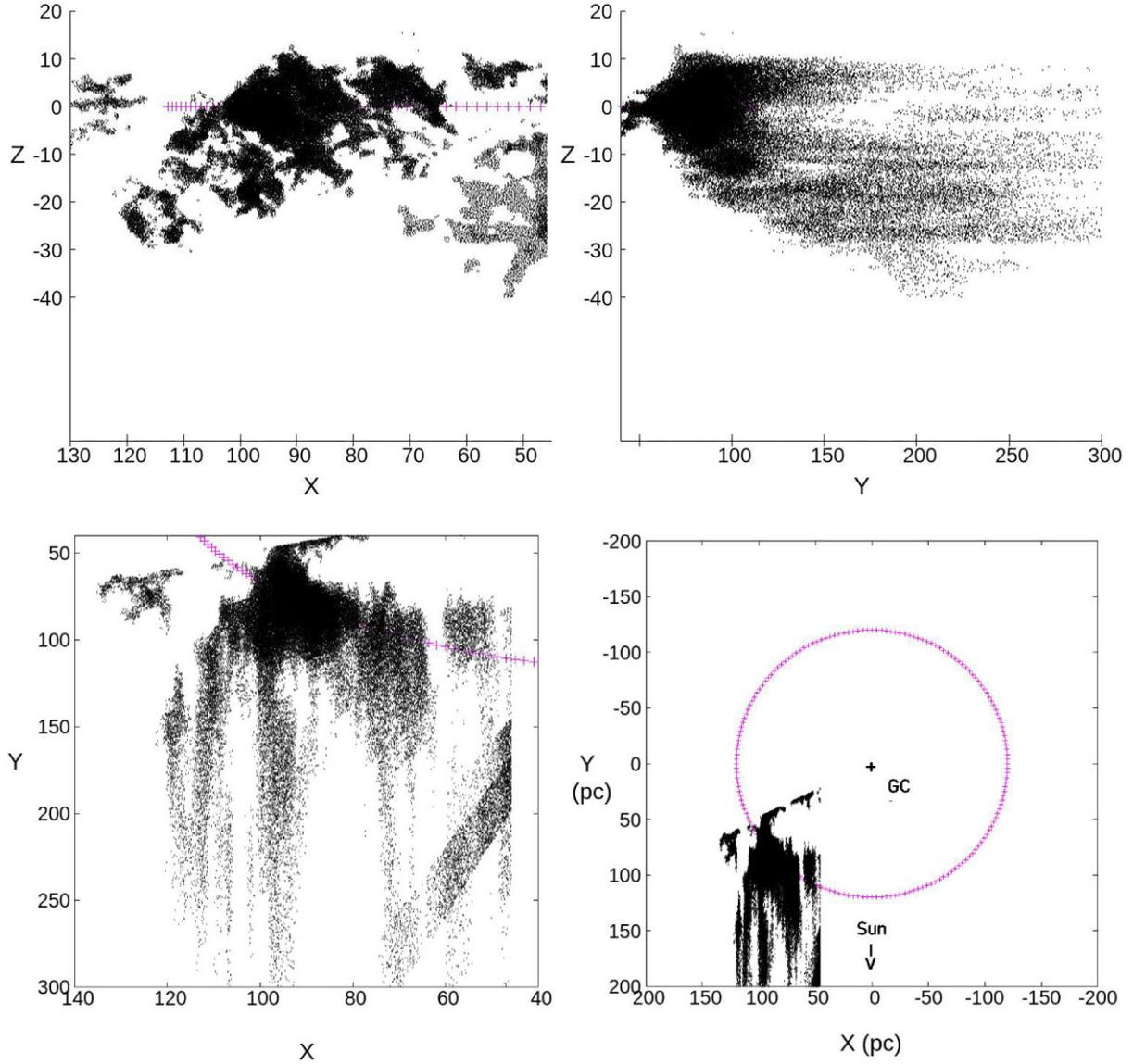


Figure 6. [Top left] Density distribution of molecular gas in Sgr-B molecular complex after applying the FOT projected on the sky ((X, Z) plane) mimicking the moment 0 map. Scale is in pc. [Top right] Same, but projected on the (Y, Z) plane, showing a side view of the complex from the east, where Y is increasing toward the Sun. Note the bow shape about the Galactic plane. [Bottom left] Same, but projected on the Galactic plane ((X, Y) plane). Magenta dots represent a ring in the galactic plane of radius 120 pc approximately representing the GC Arm I. Thus obtained 3D gas distribution composes a molecular-bow structure sheathing the Sgr-B2 SF region, making a giant cometary H II region (GCH) open toward the Sun. [Bottom right] Same, but on wider XY area including the CMZ. Sgr A* is at the centre, $(X, Y) = (0, 0)$, and the magenta circle is the 120 pc ring with GC Arm I and II (Sofue 1995). The Sun is to positive Y (downward) The outskirts of Sgr-B complex is extended for ~ 100 pc in the line-of-sight direction.

$v = v_{\text{LSR}}$ in the sense that the larger is v , the smaller becomes Y , and vice versa. As shown in the LVD in Fig. 5, the LV ridge (dashed line) representing Arm I intersects the rotation axis at $(l, v_{\text{LSR}}) \sim (0^\circ, -30 \text{ km s}^{-1})$. This implies that, if we correct for the displacement by $\delta v \sim -30 \text{ km s}^{-1}$, or replace v by $v = v_{\text{LSR}} - \delta v$, the Y -directional distance is reduced by $\sim 0.2Y$ for $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$, where the factor arising from the second term inside the square root of Fig. 3 is negligible in the present area. A similar correction is necessary, but in the opposite sense, if we adopt a different rotation velocity, so that higher V_{rot} results in larger Y distance. Thus, if we correct for the v_{LSR} displacement of $\delta v \sim -30 \text{ km s}^{-1}$ in the present data, the Y directional extent of the Sgr-B bubble may be reduced by a factor of $Q \sim 1 + \delta v/V_{\text{rot}} \sim 0.8$.

4 FORMATION MECHANISMS: MOLECULAR BOW SHOCK, COMETARY H II REGION, AND OVAL SHOCK WAVE

We have shown that the Sgr-B molecular complex has a conical-horn structure of molecular gas open toward the Sun from the Sgr-B2 cloud in the farthest end on the line of sight. In this section, we consider possible formation mechanisms of such an open bubble structure, although their local velocities are smaller than that of the Galactic rotation. We here discuss the following three cases: (a) molecular bow shock (MBS) (Sofue 2018), (b) giant cometary-H II region (GCH) (Sofue 1985, 2018, 2019; Sofue et al. 2019), and (c) oval shock wave (Sofue 1990, 2023). In the models, we assume that the complex has expanded around Sgr-B2 cloud in the flow on the

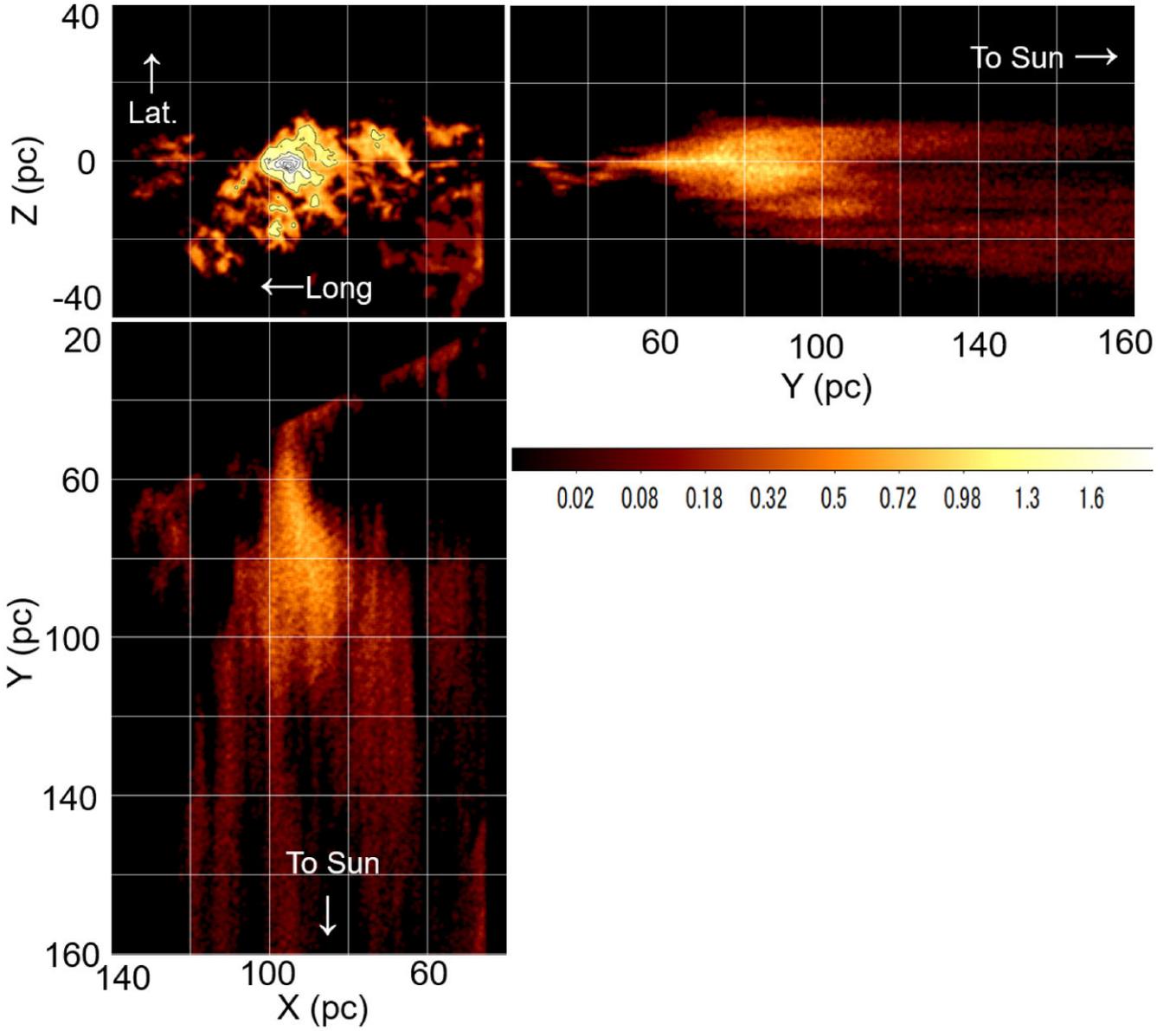


Figure 7. Same as Fig. 6, but with the same linear scaling in order to compare the X , Y , and Z extents in the same proportion. Intensity scaling is also common in all the panel in arbitrary unit. Contours in the top-left panel are drawn from 1 by interval 0.5. Note that, the Y - (v_{LSR}) directional structure is continuous from the top to tail, indicating that the entire structure is a single object.

Galactic rotation with the CMZ, which encounters the Galactic shock wave along the GC Arm I. We here also consider an expansion due to the thermal pressure of ionized hydrogen gas (H II region) heated by the central OB stars, and assume it to be on the order of the sound velocity of $v_{\text{bub}} \sim 10 \text{ km s}^{-1}$. Therefore, the bubble's age is $t_{\text{bub}} \sim r/v_{\text{bub}} \sim 2 \text{ My}$ for $r_{\text{bub}} \sim 20 \text{ pc}$. This velocity is smaller than the rotation velocity, $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$. If we take into account the expansion and differential rotation simultaneously, the front shape will be deformed in the sense that the outskirts is trailed counter-clock wise seen from the north Galactic pole due to the Coriolis force on the order of $v_{\text{bub}} \Omega r_{\text{bub}}^2 \sim 50 \text{ pc}$ for $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$ and radius $\sim 120 \text{ pc}$ (Sofue 2013), where Ω is the angular velocity of the rotation. Therefore, the deformation due to Galactic rotation is on the same order of the bubble's size. Fig. 10 illustrates the model and mechanism for the formation of the Sgr-B complex proposed in this paper based on the analysis in the previous section.

Astrophysical bow shock is a classical phenomenon in objects interacting with supersonic gas flow in the interstellar space (Dyson 1975; van Buren et al. 1990; Ogura 1995; Wilkin 1996; Arce & Goodman 2002; Reipurth et al. 2002). On the scale of star-forming site, it is observed as cometary H II regions tailing down-stream direction from the OB cluster on the head (Fukuda & Hanawa 2000; Arthur & Hoare 2006; Deharveng et al. 2015; Zhu et al. 2015; Steggle, Hoare & Pittard 2017). On galactic scale, large bow shocks at the spiral arms are observed in the tangential direction of the 4 kpc arm as associated with the SF region W43 ($l \sim 30^\circ$) and the Norma arm ($l \sim 22^\circ$) (Sofue 1985, 2019; Sofue et al. 2019). MBS are also observed in spiral galaxies such as M83, M51, and M33, and are ubiquitous in galactic shock-wave arms (Sofue 2018). Recent far-infrared images with the JWST of spiral galaxies such as M74 (NGC 628) exhibit numerous MBS and similar structures (Lee et al. 2023; Barnes et al. 2023). An MBS is a concave arc

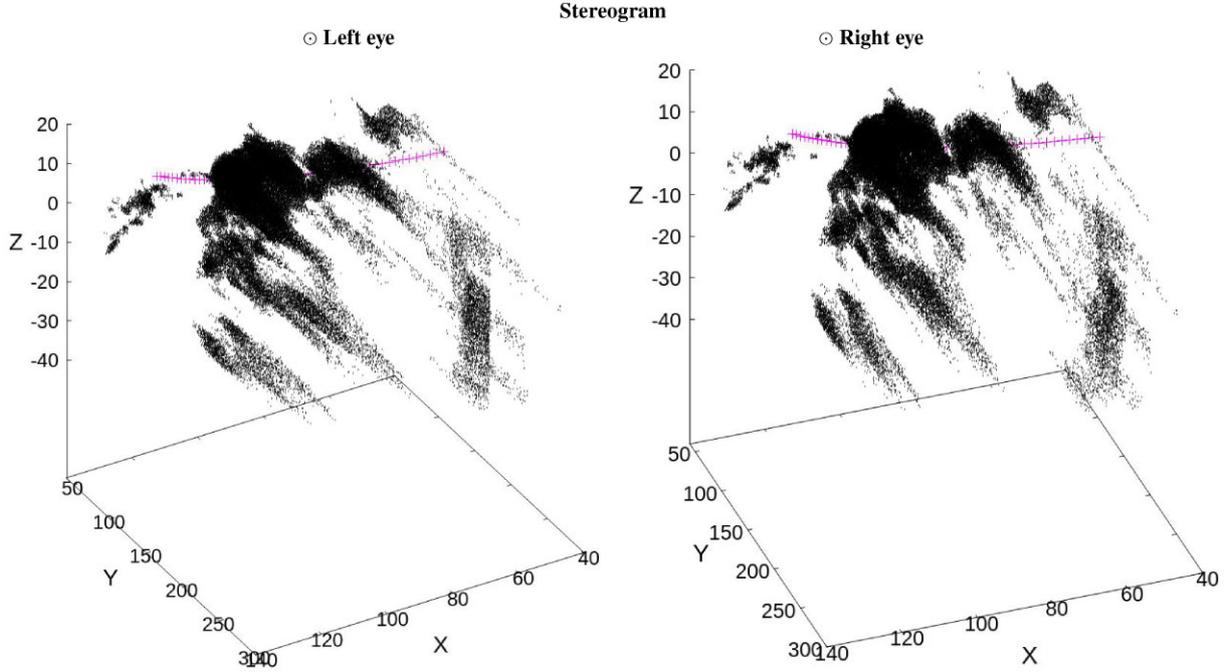


Figure 8. Oblique projections of Fig. 6 as seen from a standpoint in the north-eastern and nearer-side to the Sun. Left and right panels show projections rotated by 30° around X and 50° and 40° deg around the Z -axis. So, the two panels are seen from two directions by 10° rotated from each other, so that they can be used to obtain a stereo-graphic view of the gas distribution. Magenta dots represent a ring in the galactic plane of radius 120 pc representing the GC Arm I.

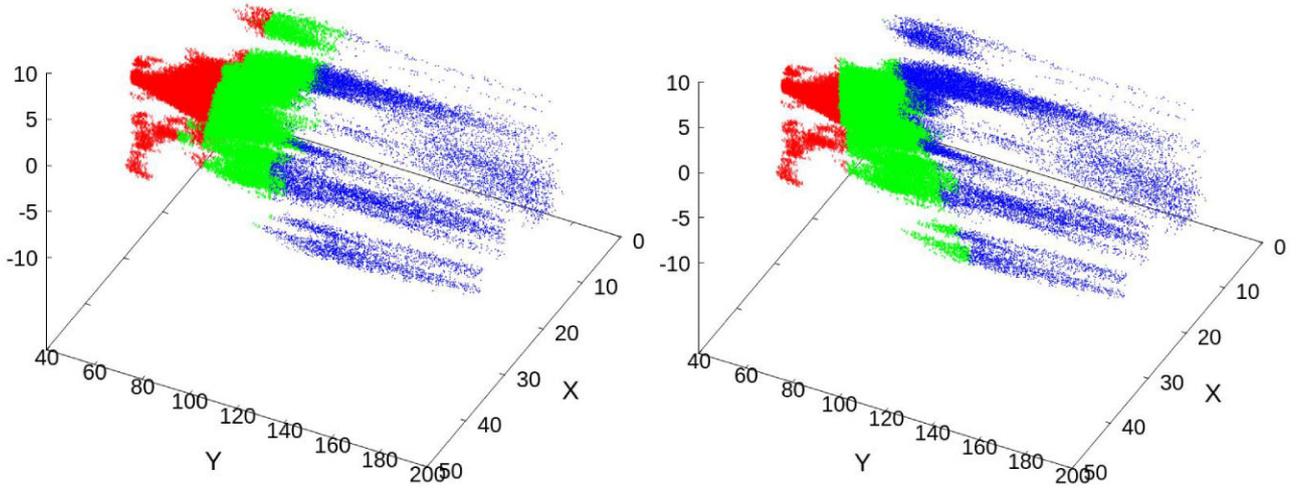


Figure 9. [Left] 3D slices of Fig. 6 projected on the sky ((X, Z) plane) at $Y < 80$ pc by red, $80 \leq Y < 110$ in green, and ≥ 110 in blue. [Right] Same, but sliced by velocity, showing large-, intermediate-, and low-velocity components in red, green, and blue, respectively.

or conical cylinder of molecular gas around an expanding H II region formed in the down-stream side with respect to the galactic shock waves. In addition to these individual phenomena, we further consider that the galactic shock-wave theory (Fujimoto 1968; Roberts 1969) applies to the arm-size structures, which include GC Arm I and II.

4.1 Molecular bow shock (MBS) in a supersonic flow

In our paper (Sofue 2018), we modelled the MBS G30.5 by applying the Wilkin (1996)'s analytical model for stellar-wind bow shock. The distance Q of a bow front from the wind source is related to the

elevation angle ϕ through

$$Q(\phi) = R_{\text{bow}} \operatorname{cosec} \phi \sqrt{3(1 - \phi \cot \phi)}. \quad (5)$$

Here, R_{bow} is the stand-off radius defined as the distance of the front on the galactic plane from the wind source, which is measured as the smallest curvature of the bow head facing the gas flow. It is related to the momentum injection rate, or the mass-loss rate of the central stars through (Vink 2022)

$$\dot{M} = 4\pi R_{\text{bow}}^2 v_{\text{CMZ}}^2 \rho_{\text{CMZ}} / v_w, \quad (6)$$

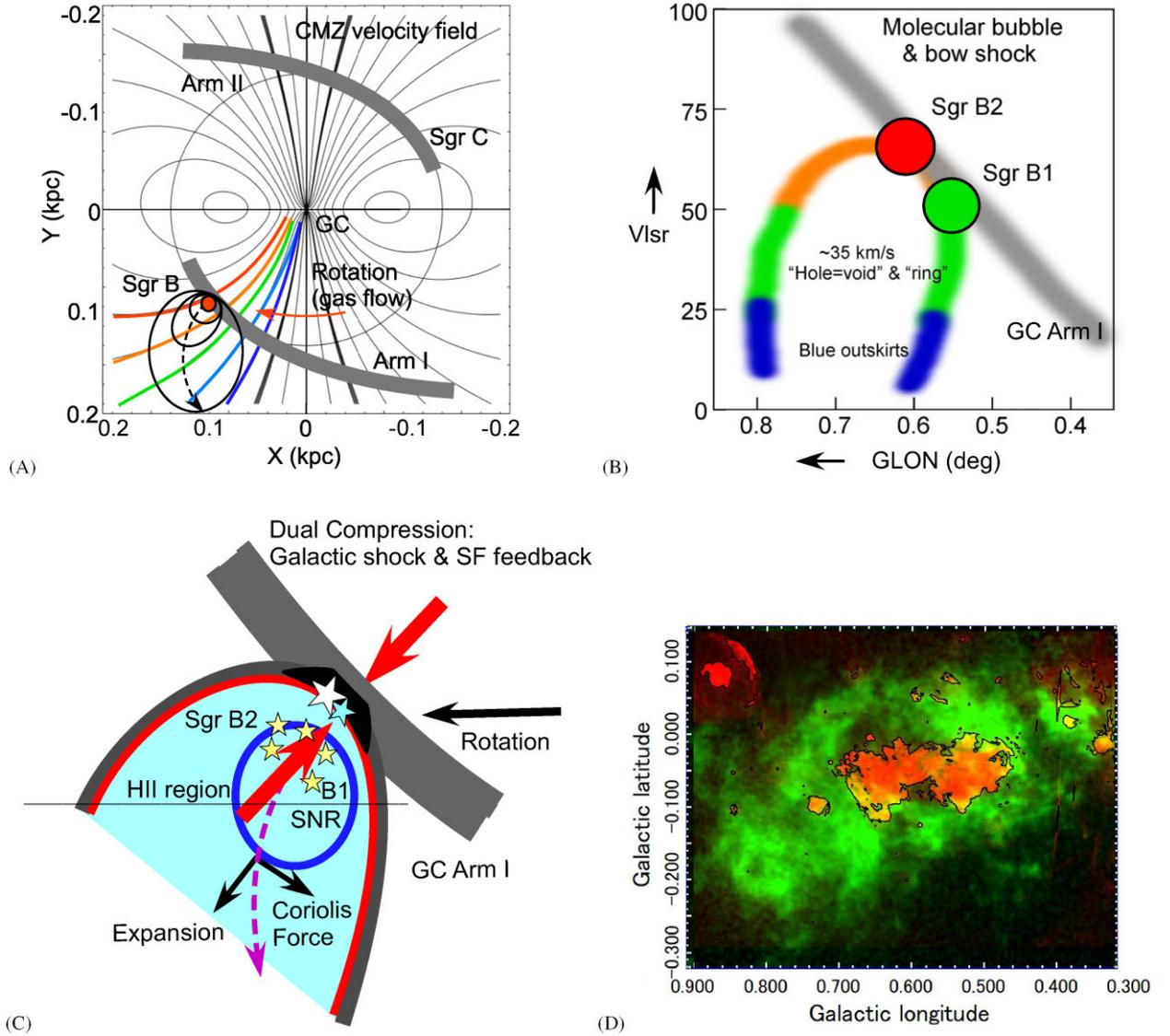


Figure 10. (A) Schematic velocity field and GC Arms I and II, and Sgr-B projected on the galactic plane. (B) Schematic longitude–velocity diagram of Sgr-B2, B2, Sgr-B molecular complex. Red colour represents $v_{\text{LSR}} \sim 60\text{--}70 \text{ km s}^{-1}$, green $\sim 30\text{--}50 \text{ km s}^{-1}$, and blue $\sim 20\text{--}30 \text{ km s}^{-1}$. The grey line is the GC Arm I on which Sgr-B1 and B2 are located. The $\sim 35 \text{ km s}^{-1}$ ‘hole and ring’ (Bally et al. 1988) is here interpreted as the slice of the bubble/cylinder at the velocity. (C) Dual-side compression (DSC) of molecular cloud by the Galactic shock wave and stellar feedback. Red arrow from right top represent compression perpendicular velocity $(V_{\text{rot}} - V_{\text{pattern}}) \sin p \sim 30 \text{ km s}^{-1}$, red arrow from left bottom. (D) MEERKAT continuum map in red with black contours at 1 mJy beam^{-1} overlaid on $^{13}\text{CO } 35 \text{ km s}^{-1}$ channel map in green (as Fig. 2 bottom left). The -molecular cloud in green approximately represents cross-section of the bow-shock cone along the straight line in panel (C), while the continuum emission is the integration along the line of sights. (Color figures online.)

where M is the mass of the central stars, R_{bow} the bow-radius at the head, v_{CMZ} is the inflow velocity of the CMZ gas, ρ_{CMZ} its density, and v_w is the wind velocity.

Figs 11 and 12 show the bow’s shape, where the wind parameter is determined to be $R_{\text{bow}} \sim 1.25 \text{ pc}$ in order to yield a reasonable fit to the observed bow of Sgr-B as in Fig. 11. The top panel of Fig. 13 shows variation of the tangential velocity as a function of the distance from the bow axis, which approximates the position–velocity diagram seen from the bow axis. The ‘V’ shape observed in the position–velocity diagrams (Figs 5) is well reproduced. The bottom panel shows the variation of the density on the bow surface. We here assumed the background gas density $\rho_{\text{CMZ}} \sim 100 \text{ H}_2 \text{ cm}^{-3} \sim 4.6 \times 10^{-22} \text{ g cm}^{-3}$, wind velocity $v_w \sim 2000 \text{ km s}^{-1}$, galactic-shock velocity of $v_{\text{CMZ}} \sim (V_{\text{rot}} -$

$V_{\text{pat}}) \sin p \sim 17 \text{ km s}^{-1}$ with $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$, $V_{\text{pat}} \sim 1/2 V_{\text{rot}}$, and $p \sim 12^\circ$ being the rotation velocity of the CMZ, pattern speed, and pitch angle of the spiral arm, respectively. Thus, we obtain $\dot{M} \sim 2.0 \times 10^{-5} M_\odot \text{ y}^{-1}$. A couple of O stars and continuously flowing molecular gas in the CMZ drive the bow structure of Sgr-B cloud.

4.2 Giant cometary H II region (GCH)

Conical structure can be modelled also by a deformed Strömgen sphere in a medium with density gradient (Sofue 2018), which yields a qualitatively similar result to full hydrodynamical simulations (Zhu et al. 2015). The Strömgen radius in a uniform ISM is related to the

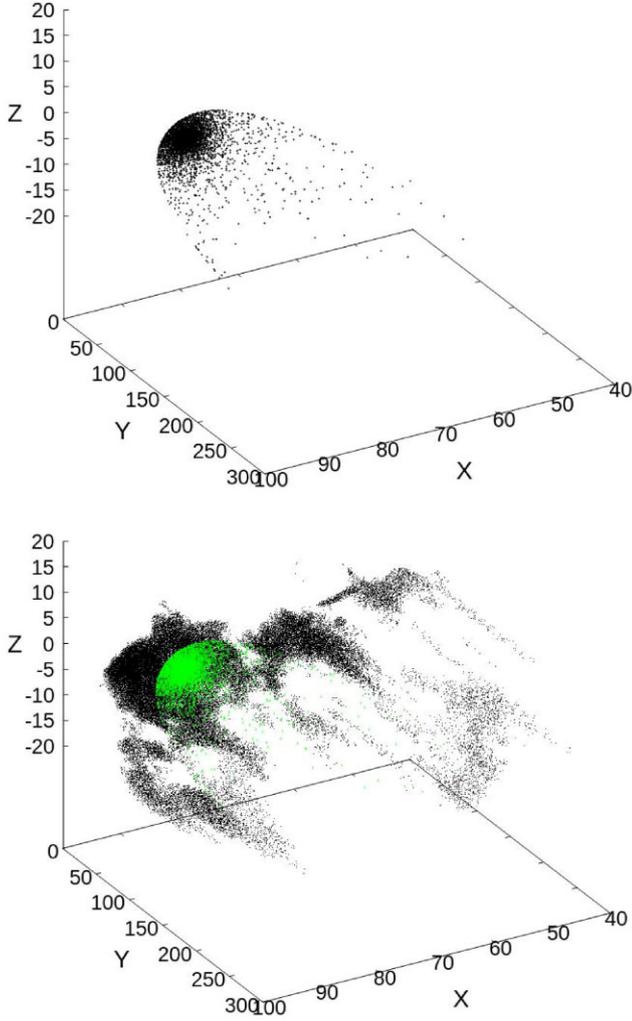


Figure 11. [Left] Bow-shock model. [Right] Overlay of 3D Sgr-B molecular complex and bow-shock model.

luminosity of the central OB stars by

$$\mathcal{N}_{\text{UV}} \simeq \frac{4\pi}{3} R_{\text{HII}}^3 n_i n_e a_r, \quad (7)$$

where \mathcal{N}_{UV} (s^{-1}) is the UV photon number radiated by the OB stars per time, $a_r \sim 4 \times 10^{-13} \text{ cm}^{-3} \text{ s}^{-1}$ is the recombination rate, n_i and n_e are the ion and electron densities, respectively.

We here assume that the equation holds in each small solid angle at any direction in which the density is assumed to be constant. By neglecting dynamical motion of gas, this approximation gives a qualitative shape of the expanding H II-region front.

Assuming that the ion and electron densities are related to the neutral gas density n through $n_e \sim n_i \sim n(T_n/T_e)$, where T_n and T_e are the temperatures of neutral and H II gas, we define the representative radius R_{HII} as the equilibrium radius of a spherical H II region in uniform gas with density n by

$$R_{\text{HII}} \simeq \left[\frac{3\mathcal{N}_{\text{UV}}}{4\pi a_r n^2} (T_e/T_n)^2 \right]^{1/3}. \quad (8)$$

Rewriting $\mathcal{N}_{\text{UV}} \sim \mathcal{L}/h\nu$ with \mathcal{L} and $h\nu$ being the luminosity of the central OB stars and UV photon energy over 912 Å, respectively, we

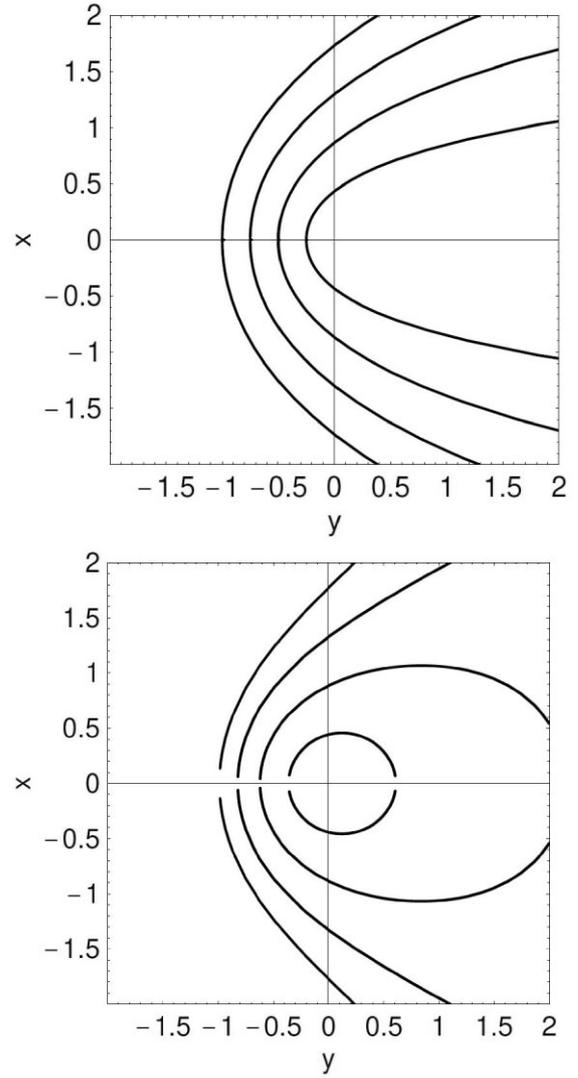


Figure 12. [Top] Bow-shock front by supersonic flow encountering with stellar wind given by equation 5, where x and y are measured from the star in unit of the stand-off radius R_{bow} (Wilkin 1996). [Bottom] Expanding H II-region front in a density gradient proposed for the molecular bow shock and cometary H II regions (Sofue 2018). The scale is arbitrary.

have

$$R_{\text{HII}} \sim 96 \left(\frac{\mathcal{L}}{10^4 L_{\odot}} \right)^{1/3} \left(\frac{n}{10^2 \text{ H cm}^{-3}} \right)^{-2/3} \left(\frac{T_n}{20 \text{ K}} \right)^{-2/3} \times \left(\frac{T_e}{10^4 \text{ K}} \right)^{2/3} (\text{pc}). \quad (9)$$

The front shape varies according to the parameter $R_{\text{HII}} \propto \mathcal{L}^{1/3} n^{-2/3}$, and attains an elliptical shape elongated in the lower-density direction (Sofue 2018). On the other hand, given the gas density, $n \sim 10^3 \text{ H}_2 \text{ cm}^{-3}$, the UV luminosity \mathcal{L} of exciting OB stars can be estimated for the measured $R_{\text{HII}} \sim 20 \text{ pc}$ as $\mathcal{L} \sim 10^4 L_{\odot}$. The bottom panel of Fig. 12 shows an example of such deformed H II front proposed for the cometary H II regions in M83 (Sofue 2018).

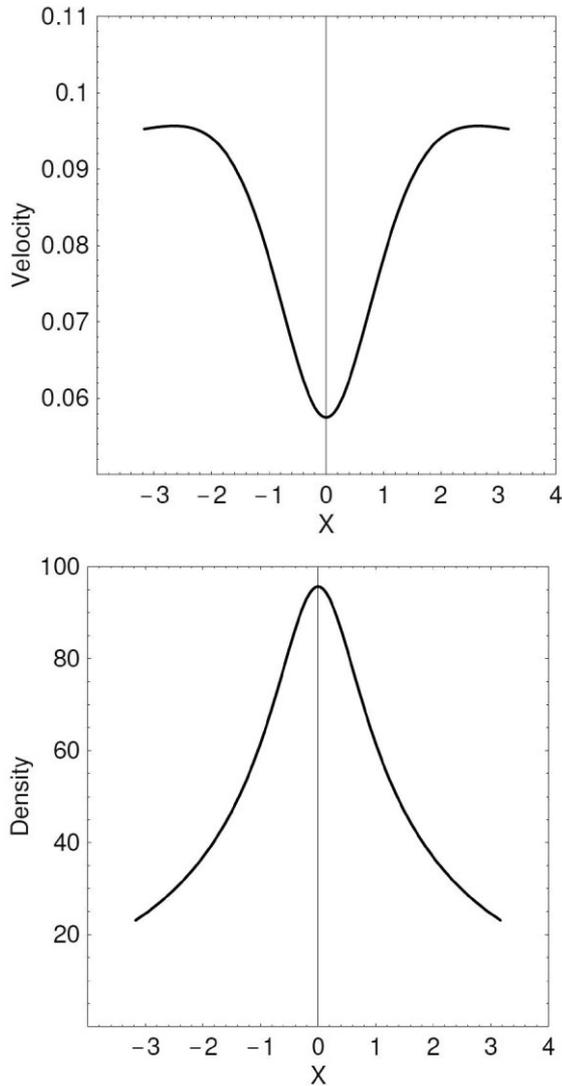


Figure 13. [Top] Tangential velocity of the flow along the bow-shock plotted against the distance from the bow axis corresponding to Fig. 12 following Wilkin (1996). Vertical axis is arbitrary. [Middle] Corresponding surface-gas density against. Vertical axis scale is arbitrary.

4.3 Oval shock front driven by H II region and/or SNR

The third mechanism is a shock wave driven by supersonic expansion of the H II region and/or supernova remnant in Sgr-B2 and B1. In order to examine this model, we apply the Sakashita’s method in order to trace the propagation of an adiabatic-shock wave in an inhomogeneous medium originally developed to trace the propagation of a shock-wave from a galactic nucleus through the disc (Sakashita 1971). The method has been applied to various galactic objects including an interstellar shock wave from a molecular cloud. We mimic here the original dense cloud by an elongated density concentration in a uniform medium as investigated in our recent paper (Sofue 2023).

In this model, energy on the order of $\sim 10^{50}$ erg is released at a point on the oblique edge of the cloud, and the propagation of the shock front is traced using Sakashita’s method. Fig. 14 shows an example of such oblique shock front projected on the sky and corresponding position–velocity diagram (right panels) as compared with the intensity distribution and latitude–velocity diagram of the Sgr-B complex (top left and bottom right, respectively).

The shock front expands spherically in the initial stage, but after a while it is deformed to attain an oval structure, and expands faster in the direction to decreasing density. The top-right panel of Fig. 14 shows such elongated oval projected on the sky, while the projected shape is nearly spherical because the line of sight is nearly parallel the oval’s major axis in this model.

The bottom-right panel shows a position–velocity diagram across this front as seen from the major axis (observer), and is compared with the observation (left panel). The shock front draws a lopsided oval, mimicking the observed LVD.

We have thus shown that an explosive event at the cloud edge can explain the lopsided bubble structure as well as the position–velocity behaviour. In these calculations, however, we did not take into account the rotation of the ambient medium, which would deform the bubble shapes on the sky as well as in the LVD. Although the present models may explain the observed shapes qualitatively, full dynamical simulations are necessary to quantitatively reproduce the observed properties.

5 DISCUSSION

As described in Section 1, there have been three models so far proposed for explaining the ‘hole’ and ‘ring’ of the Sgr-B molecular complex, which are (i) the local perturbation model (Bally et al. 1988), (ii) expanding-bubble and stellar feedback model (Sofue 1990), and (iii) cloud-collision model (Hasegawa et al. 1994; Sato et al. 2000; Enokiyama & Fukui 2022). The purpose of this paper was to explore the bubble/shell model (ii) in the light of recent CO-line and radio-continuum data. We explained the complex by an MBS (molecular bow shock) induced by the galactic shock wave in the CMZ, which produced the a horn-like bubble structure assisted by the energy feedback of star formation in Sgr-B1 and B2 associated with a GCH (giant cometary H II region). We have shown that the model can explain various properties observed in the complex, which we summarize below. We summarize the present model in analogy to the bubbles and bow-shock structures as observed in M74 and M83 in Fig. 15.

5.1 Galactic shock waves, fins, and voids

The galactic shock wave induced by the density waves of spiral arms (Fujimoto 1968; Roberts 1969; Roberts, Huntley & van Albada 1979; Wada, Baba & Saitoh 2011; Baba, Morokuma-Matsui & Saitoh 2017) and bar potential (Sorensen, Matsuda & Fujimoto 1976; Sormani, Binney & Magorrian 2015) are the well-studied and established phenomenon in the gas dynamics of galactic discs, which also applies to the Galactic centre (Sormani et al. 2020). High-resolution numerical simulations show that fin structures are bifurcated from the spiral shock waves and compose void features (Baba et al. 2017). Recent far infrared imaging with the JWST of nearby spiral galaxies revealed numerous ‘voids’ in the dusty spiral arms, which are generally curved toward the up-stream direction (Lee et al. 2023; Barnes et al. 2023). We point out that the voids are far-infrared views of the MBS (molecular bow shocks) sheathing GCH (giant cometary H II regions) (Sofue 2018). We argue that the MBS + GCH structure is ubiquitous also in the Milky Way (Sofue 1985; Sofue et al. 2019) and the CMZ as observed as Sgr-B complex here. The result of this paper lies on the extension of these general views of the molecular gas distribution in spiral arms and central bar-shocked regions of galaxies.

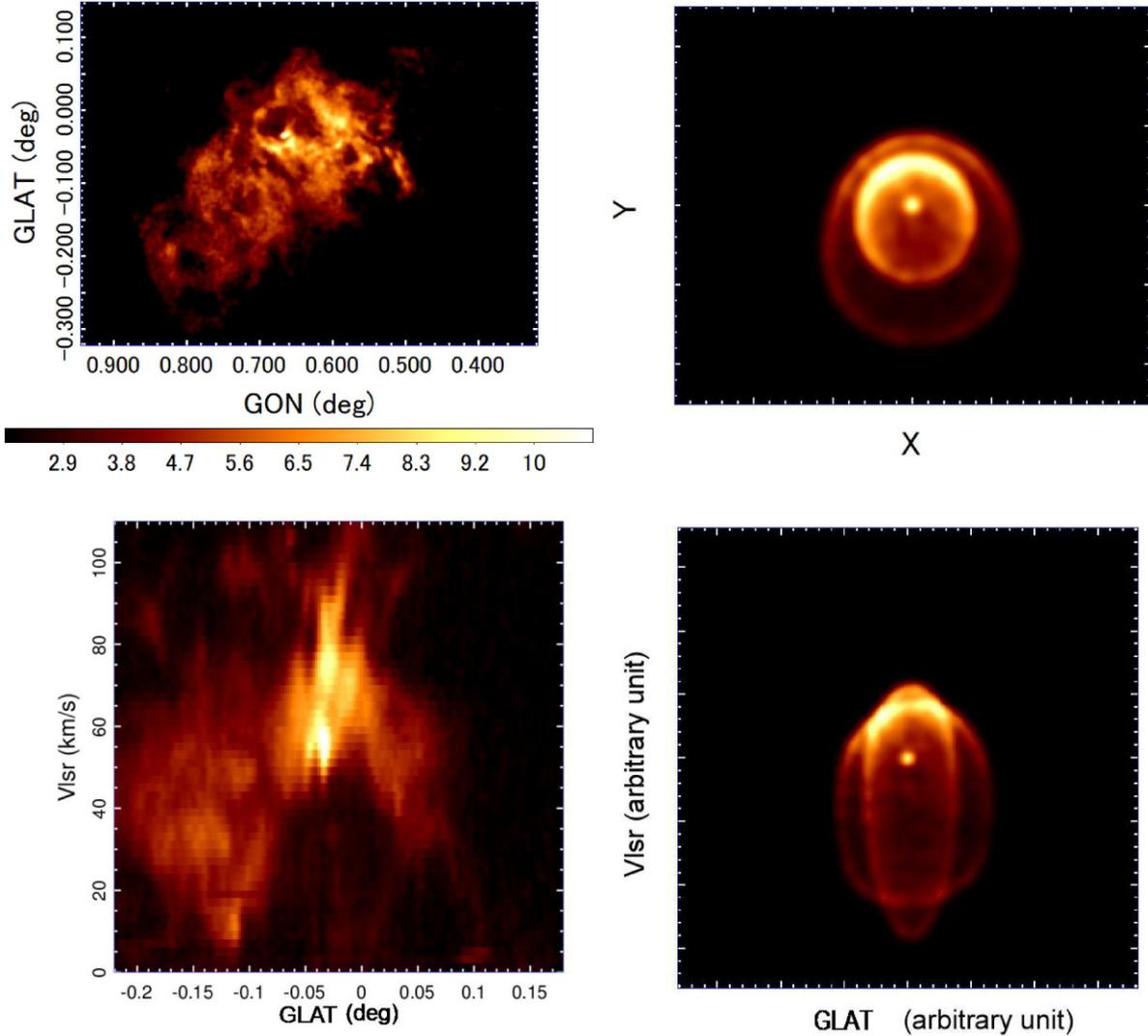


Figure 14. [Top left] Sgr-B molecular complex in the ^{13}CO -line emission. [Top right] Calculated shock-wave fronts projected on the (X, Y) plane due to an explosion at the cloud's edge. Fronts at two epochs at $t = 1$ and 2 time units from the explosion are presented in order to show the evolution. The shell expands toward the low-density side (toward bottom), attaining a lopsided open cone, and mimics the molecular bubble around Sgr-B. Scales are arbitrary and the time is normalized by r_0/v_0 with r_0 and v_0 being the unit radius and expansion velocity in a uniform medium with unit density, respectively. [Bottom left] Observed latitude–velocity diagram across Sgr-B2. [Bottom right] Calculated position–velocity diagrams of the above shock fronts seen from the opening end of the cone. This model demonstrates that the bubble and V-shaped PVD in Sgr-B can be produced also by such a lopsided explosion, which suggests that the real Sgr-B may be a hybrid of asymmetric bubbles formed by various mechanisms.

5.2 ‘V’-shaped position–velocity diagrams around stagnation

The linear alignment of the Sgr-B1 and B2 clouds on the LVD (Fig. 5) indicates that the two clouds are located on the GC Arm I and are rotating circularly in the Galactic disc. It is remarked that a number of compact H II regions are associated with Sgr-B1 and B2 as observed in radio recombination lines, which lie on the same LV ridge (Mehring et al. 1992, 1993). This means that Sgr-B2 cloud composes a part of the main structure of Arm I in regular Galactic rotation, but it is not a flying-by object, as the collision model postulates.

The ‘V’-, or ‘ Λ ’, shaped LVD near Sgr-B2 (Fig. 5) is explained by intersection of the main ridge of GC Arm I and the horn-like bubble of the Sgr-B complex, where the main LV ridge (Fig. 5) nests the H II regions Sgr-B2 and B1. The H II regions are rotating on the Arm, exhibiting the LSR-velocity gradient

$dv_{\text{LSR}}/dl \simeq 150 \text{ km s}^{-1}/\text{deg} \simeq 11 \text{ km s}^{-1} \text{ pc}^{-1}$. The western wing of the Λ is, therefore, the main ridge of the rotating Arm, but not due to local disturbance from outside. The molecular bubble shows up as a bow-shaped arc in the LVD in touch with the GC Arm main ridge (Figs 5 and 10), where the western wing composes a fainter Λ , bifurcating from the main Arm toward decreasing velocity. The LV behaviour is thus naturally explained by the galactic rotation and extended bubble in the rotating disc.

The latitude–velocity diagrams (BVD) across Sgr-B2 also exhibit V shape as shown in Figs 4, 5, and 14. It is interesting to point out that the width of the ‘ Λ ’ in the BVD is $\sim 0^\circ.15$ (20 pc), about a half of the width in LVD of $\sim 0^\circ.3$ (40 pc). This flattened conical shape in the (l, b, v_{LSR}) space can be naturally understood as due to a bow shock by the galactic shock wave flattened along the disc plane.

We mention that the V-shaped bow structure is also clearly seen in the HCN ($J = 4-3$) line, which is the highest-density tracer among

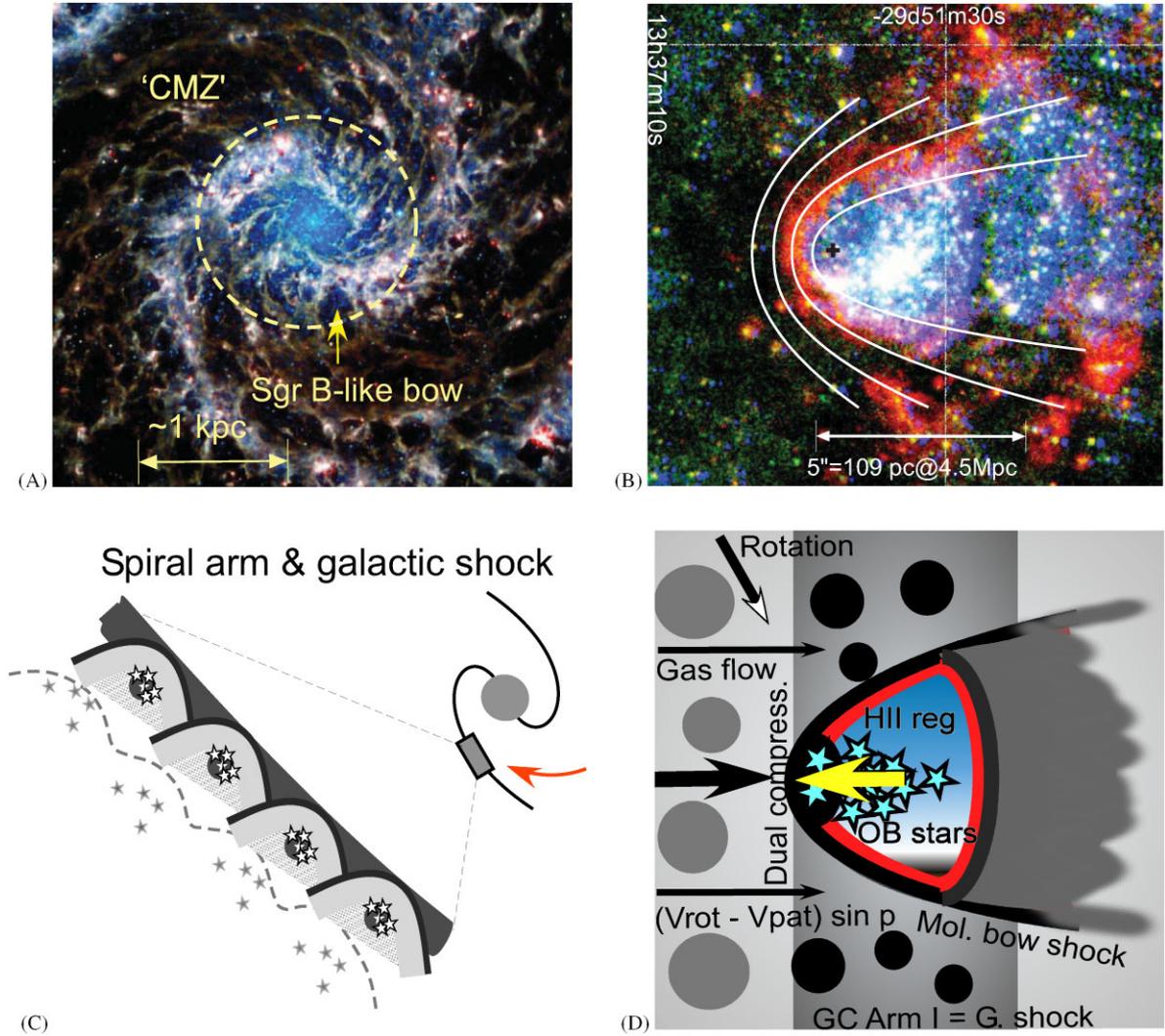


Figure 15. (A) Molecular bow shock similar to Sgr-B molecular complex in the mid-infrared image of the spiral galaxy M74 (NGC 628) taken with the JWST (Lee et al. 2023; Barnes et al. 2023). Numerous MBS are expanding toward the down-stream sides of the spiral arms. (B) Giant cometary HII region (GCH) sheathed by an MBS in the optical image of M83 with the *HST* (McQuinn et al. 2017). The front is fitted by a bow-shock as indicated by the white lines (Sofue 2018). (C) Illustration of a chain of MBS+GCH along a spiral galactic shock wave. (D) Illustration of the MBS + GCH model for Sgr-B molecular complex. Dual-side compression (DSC) of the stagnated molecular cloud triggers the star formation. Velocity vectors (arrows) are relative to the rotating arm, where V_{rot} , V_{pat} , and p are the rotation velocity, pattern speed and arm’s pitch angle, respectively. (Color figures online.)

the lines ever used for the CMZ (Tanaka et al. 2018), as shown in panel (D) of Fig. 5. Moreover, the stagnation point (cap) is saturated in the line, even producing deep absorption against the continuum emission of Sgr-B2. A similar bow feature with absorption at Sgr-B2 is also recognized in the CS line (Tsuboi, Handa & Ukita 1999). The fact indicates that the gas density attains a steep maximum at the stagnation point associated with the active star formation of Sgr-B2, where the critical density is $\gtrsim 10^7 \text{ H}_2 \text{ cm}^{-3}$, while the effective density would be one or two orders of magnitude lower due to the photon trapping effect (Tanaka et al. 2018).

All the models for Sgr-B complex presented in Section 4 indicate similar behaviours of flow speed in the position–velocity diagram. In the bow-shock model, the supersonic flow is stagnated at the bow head, where the velocity is zero with respect to the shock wave, and increases along the bow surface with the distance from the bow axis (Wilkin 1996). As shown in Fig. 13, the theoretical calculation of the tangent velocity along the bow as a function of the distance

from the bow axis well explains the observed V shape in Fig. 5. We also point out that the theoretical density variation along the bow, as shown in the bottom panel of Fig. 13, is well observed in the BVD across Sgr-B2 (Fig. 5). In the expansion models due to H II regions, a similar V shape appears according to the lopsided expansion at different velocities depending on the directions from the centre (Fig. 14).

Thus, in general, a V-shaped PVD is associated with stagnated flow of gas, which is observed in the galactic shock waves, in bow shocks around a star, around an H II region, around a cloud, and in a cloud under collision with another cloud. Lopsided expanding shells around H II regions in non-uniform medium also exhibit partially elliptical PVD, mimicking V shape. Therefore, the V-shaped PVD is a general astrophysical phenomenon in the ISM. We mention that a similar property is observed in the cloud-collision, where the flow pattern around the colliding body is identical to that in general formed around the bow shock.

5.3 Complicated velocity dispersion in the Sgr-B2 cloud

The moment 2 map in Fig. 1 indicates systematically lower velocity dispersion at $\sigma_v \sim 10$ toward the Sgr-B2 cloud over $0^\circ.07$ (10 pc) diameter than the surrounding cloud with $\sigma_v \sim 18 \text{ km s}^{-1}$. This fact is opposite to that expected for a cloud collision which induces stronger turbulence and velocity dispersion between the two clouds. This fact is in favor of the expanding-bubble model, where the central high-density Sgr-B2 cloud composes the major structure of the GC Arm I, and the expelled clouds from the Arm compose the expanding bubble with higher velocity dispersion.

On the other hand, the very central clump of diameter of $0^\circ.02$ (3 pc) at $(l, b) = (0^\circ.66, -0^\circ.03)$ near the centre of sgr-B2 cloud exhibits an anomalously high velocity dispersion of $\delta v \sim 40 \text{ km s}^{-1}$ around the centre velocity $v_{\text{LSR}} \sim +72 \text{ km s}^{-1}$ (Fig. 5). Such a large velocity width cannot be explained neither by the expansion of an H II region, which is on the order of $\sim 10 \text{ km s}^{-1}$, nor by the cloud collision at colliding speed $\sim 20\text{--}30 \text{ km s}^{-1}$. This compact cloud is near to the high-velocity clump no. 138, $(l, b, v_{\text{LSR}}) = (0^\circ.671, -0^\circ.032, +94.0 \text{ km s}^{-1})$ listed as a ‘wing’ in the catalogue by Oka et al. (2022). However, the centre velocity is significantly different, so that their physical relation is not clear at the moment.

5.4 Dual-side compression (DSC) of molecular cloud and high efficiency of star formation

MBS is composed of low-temperature and high-density molecular gas. On the other hand, GCH is high-temperature and low-density H II gas expanded around OB clusters. Both the MBS and GCH show a similar morphology as the GCH is sheathed by MBS after their interaction and mutual deformation. The expansion of GCH is suppressed by MBS on the up-stream side, while it becomes more free in the tail. The side wall of MBS also suppresses the H II expansion, and guides the H II gas into a sharper cone.

During their interaction, the GCH and MBS act to trigger star formation in such a way that GCH compresses the stacked molecular gas at the bow head from inside, and MBS compresses the molecular clouds from outside (Fig. 10). The supersonic flow by the spiral arm’s gravitational potential convey and supply molecular clouds to the bow head. Thus, the bow head becomes an efficient SF site by the ‘dual-side dynamical compression’ (DSC) by the MBS and GCH.

The frequency for one cloud to experience the DSC is equal to the frequency for a cloud to encounter the spiral arms in the CMZ, which is given by $f \sim (V_{\text{rot}} - V_{\text{pat}})/(\pi R)$, where V_{rot} is the rotation velocity, $V_{\text{pat}} = \Omega_p R$ the pattern velocity, and Ω_p is the angular speed of the pattern. If we assume that $R \sim 100 \text{ pc}$, $V_{\text{rot}} \sim 150 \text{ km s}^{-1}$ and the pattern’s angular velocity is equal to that in the Galactic disc, $\Omega_p \sim 20 \text{ km s}^{-1} \text{ kpc}^{-1}$, yielding $V_{\text{pat}} \sim 40 \text{ km s}^{-1}$, we obtain $f \sim 4 \times 10^{-7} \text{ yr}^{-1}$.

5.5 MBS in extra-galactic CMZ

The bow-shock feature is ubiquitous not only in the Galaxy, but also in spiral galaxies. In Fig. 15, we show a bow shock calculated in order to simulate a giant MSB and GCH observed in M83, which are well fitted by fronts with bow radii $R_{\text{bow}} \sim 25\text{--}60 \text{ pc}$ (Sofue 2018). Similar bow shocks are observed also in the recent JWST images of nearby spiral galaxies such as M74 (Lee et al. 2023; Barnes et al. 2023) as shown in the right panel of Fig. 15. So, we may conclude that Sgr-B molecular complex is one of the typical bow-shock phenomena in disc galaxies.

5.6 Evolution of Sgr-B molecular complex

Finally, we propose a possible scenario of evolution of the Sgr-B1 and B2 clouds based on the observational facts presented in this paper. We consider that the CMZ is an ordinary gaseous disc rotating in the Galactic potential, which is superposed by a spiral density wave or a bar potential. The orbits of the ISM and the molecular clouds are disturbed to cause the galactic shock wave in the oval potential well. Molecular gas and clouds entering the shock wave are strongly compressed to form stars. Each molecular cloud is compressed not only by the increase of the pressure of the surrounding gas, but also by the stagnation at the bow-shock head, as well as by the counter-compression by the thermal pressure of the H II region expanding from the newly born OB stars. This view is ubiquitous in the discs of the Milky Way and in spiral galaxies, and therefore, the Sgr-B region can be regarded to be a normal star forming cloud.

Fig. 16 illustrates the evolution of the Sgr-B1/B2 complex. A molecular cloud orbiting around the GC enters the GC Arm I, where star formation is triggered by the galactic shock compression, and produces Sgr-B1 SF region. Then, a neighbouring bigger cloud enters the shock and is compressed to form denser cloud followed by star, which creates Sgr-B2, where the feedback of the prior SF in Sgr-B1 enhanced the SF by the pressured of the expanding H II region. The subsequent gas flow from behind causes a bow shock headed by the Sgr-B2 cloud, whose outskirts in the down-stream side composes the extended molecular bubble. As the SF region and molecular bubble around Sgr-B2 evolve, Sgr-B1 is merged by the molecular bow of Sgr-B2, which is observed at the present time. This sequential process of the cloud compression, bow shock, star formation, SF feedback, and dual compression of the bow head is a general gas dynamical mechanism in a spiral arm of a galactic disc. This mechanism will therefore also explain the origin of the star-formation belt in the entire CMZ disc (Sofue 2022). As the molecular complex expands, it is trailed toward the downstream by the Galactic rotation, attaining an elongated structure toward the Sun as well as to the east on the sky, which is indeed observed as the lopsidedness of the complex as shown in Fig. 2. The scenario would be more general and may be applied to the network of deformed dust voids containing SF regions as widely observed in spiral galaxies (Lee et al. 2023; Barnes et al. 2023) (Fig. 15). However, the vertical displacement of the whole complex by $\sim 30 \text{ pc}$ toward south from the Galactic plane remains an unanswered question.

5.7 Comparison with other models

We comment on the other two models so far proposed for the Sgr-B molecular complex.

Bally et al. (1988) interpreted the ‘hole and ring’ at $v_{\text{LSR}} \sim 30 \text{ km s}^{-1}$ as due to gravitational perturbation by the massive Sgr-B2 cloud at $\sim 60 \text{ km s}^{-1}$. This idea is interesting in the sense that we may be watching a contraction phase of a giant cloud by the self gravity, because the central region contracts faster than outskirts due to shorter Jeans time. Difficulty in this model would be the difference in the systemic-velocity of $\delta v_{\text{LSR}} \sim 30\text{--}40 \text{ km s}^{-1}$. If we consider the Galactic rotation, the mutual distance would be too far for gravitational attraction between the core and ring. In this model, we interpret the hole as a feature due to velocity slicing at $\sim 30\text{--}40 \text{ km s}^{-1}$ of a conical cylinder continuously distributed in the (l, v, v_{LSR}) space.

In the cloud-collision model (Enokiya & Fukui 2022), the $70\text{--}100 \text{ km s}^{-1}$ cloud associated with Sgr-B2 and the $40\text{--}60 \text{ km s}^{-1}$ cloud are assumed to be two distinct unbound bodies orbiting on a straight line

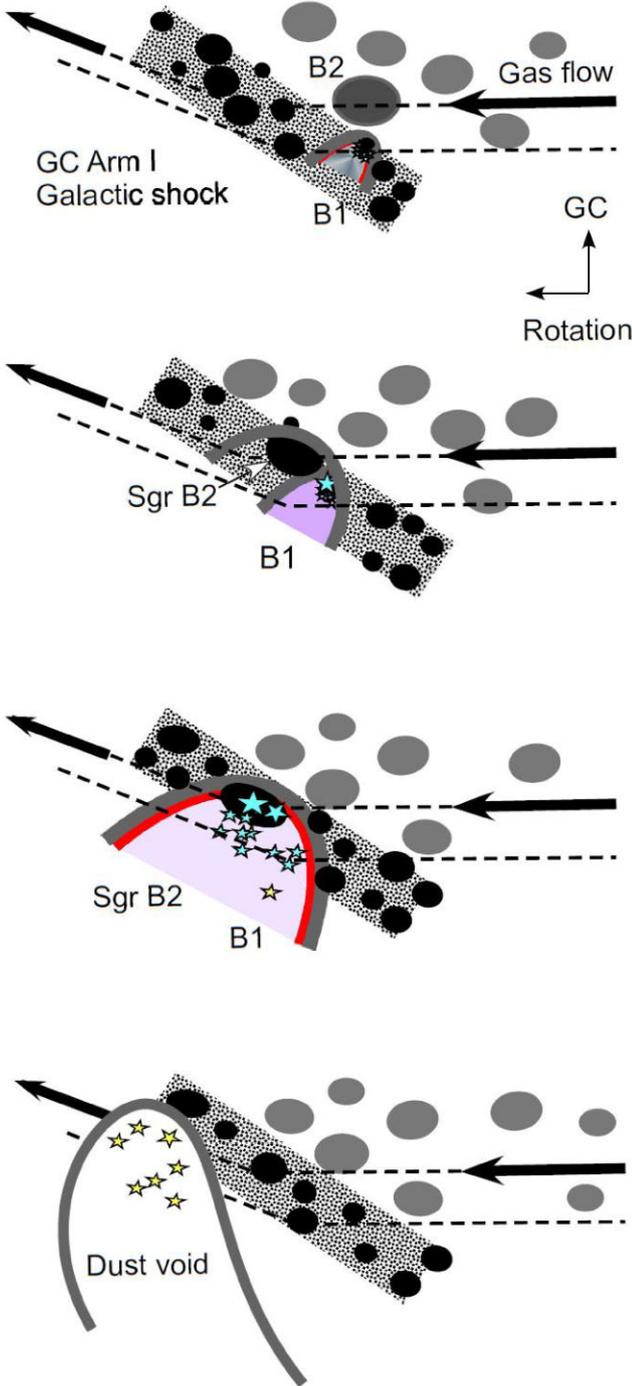


Figure 16. Evolution of the Sgr-B1 and B2 clouds in the GC Arm I, which is a galactic shock wave in the central oval potential. This view is ubiquitous in the discs of the Milky Way and spiral galaxies. This scenario will apply to the network of deformed dust voids containing SF regions as observed generally in spiral galaxies such as M74 (Lee et al. 2023; Barnes et al. 2023) (Fig. 15).

(extremely hyperbolic orbits) with zero angular momentum. In the circumstance of CMZ, the mean-free path and collision frequency of clouds of radius ~ 10 pc randomly moving at $v_c \sim 30$ km s $^{-1}$ are estimated to be $L_{\text{mfp}} \sim 400$ pc and $t_c \sim 10$ Myr, respectively. Since the 70-km s $^{-1}$ cloud (Sgr-B2) composes the main LVD ridge of GC Arm I rotating with the disc of CMZ, the 40-km s $^{-1}$ cloud would be

the passerby cloud. If so, the 40-km s $^{-1}$ cloud had to come from the far side of Sgr-B2, flying in the counter sense to the Galactic rotation.

A concern about this model is the assumed zero angular momentum. In order to prove such a special (straight) orbit in order for a head-on collision of two unbound objects, determination of the six orbital elements is the inevitably necessary condition, which, however, seems to have been not obtained.

Another concern is that the collision model ignores the rapid Galactic rotation in the deep gravitational potential of the Galactic Centre, which prohibits such a straight or hyperbolic orbit. In the rotating system, the angular velocity $\Omega = V_{\text{rot}}/R$ yields the Coriolis force (acceleration) $a_{\text{cor}} = -2\Omega \times v_c$ on the object moving at velocity v_c with respect to the rotating coordinates. The object therefore attains the epicyclic motion, drawing a local circle of radius $R_{\text{epi}} = v_c/(2\Omega)$ with a period of $P_{\text{epi}} = \pi/\Omega$. For the 40-km s $^{-1}$ cloud having $V_{\text{rot}} \sim 150$ km s $^{-1}$, $R \sim 100$ pc, and $v_c \sim 30$ km s $^{-1}$, we obtain $R_{\text{epi}} \sim 10$ pc and $P_{\text{epi}} \sim 2$ Myr. Namely, the 40-km s $^{-1}$ cloud is orbiting around the epicyclic centre 10 pc displaced from its present position. Since Sgr-B2 is rotating at the same angular speed (Ω) around the GC, the 40-km s $^{-1}$ cloud encounters/passes Sgr-B2 every $P_{\text{epi}} \sim 2$ My during its epicyclic oscillation.

Such difficulty does not exist in our bubble model, which attributes the whole Sgr-B complex to a large single system rotating with the disc of CMZ. The continuous distribution of gas is indeed confirmed in the channel maps (Fig. 3), position–velocity diagrams (Fig. 5), and in the FOT maps (Fig. 7)

6 SUMMARY

We revisited the expanding-bubble model for the Sgr-B molecular complex (Sofue 1990) based on the analysis of high-resolution CO-line cube data of the Galactic Centre observed with the Nobeyama 45-m radio telescope by Tokuyama et al. (2019). Considering the rotation of the CMZ at $V_{\text{rot}} \sim 150$ km s $^{-1}$ and applying the FOT method to the CO-line cube, we derived the 3D structure of Sgr-B molecular complex. The 3D molecular-gas distribution exhibited a conical-horn structure with the densest Sgr-B2 cloud at $v_{\text{LSR}} \sim 70$ km s $^{-1}$ in the farthest end followed by the conical bubble at $v_{\text{LSR}} \sim 60$ – 20 km s $^{-1}$ open toward the Sun. The entire Sgr-B complex is elongated in the direction of the Sun with full 3D sizes of $\Delta X \times \Delta Y \times \Delta Z \sim 40$ pc \times 100 pc \times 20 pc, where Z is the axis toward increasing longitude, Y is the axis toward the Sun, and Z toward the north.

We proposed a formation mechanism of the derived 3D structure of Sgr-B complex as due to a molecular bow shock associated with a giant cometary H II region in the galactic shock wave in the GC Arm I of the CMZ. We further argued that dual-side compression (DSC) of the molecular cloud by the Galactic shock wave and expanding H II region acts to enhance the star formation in the GC Arm I, resulting in the active SF regions Sgr-B1 and B2, whose feedback produced the conical-horn shape of the molecular complex. We conclude that the star formation in the GC is triggered by the Galactic-scale shock waves propagating in the dense CMZ as well as by the feedback through expansion of H II regions.

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CONFLICT OF INTEREST

The author declares that there is no conflict of interest.

DATA AVAILABILITY

The CO data were taken from <https://www.nro.nao.ac.jp/~nro45mr/html/results/data.html>. The MeerKAT radio data were downloaded from <https://archive-gw-1.kat.ac.za/public/repository/10.48479/fyst-hj47/index.html>. The optical image of M83 was reproduced from the web sites of STSci at <http://www.stsci.edu/hst/wfc3/> and NASA at <https://apod.nasa.gov/apod/>. The infrared image of M74 was taken from url: https://www.esa.int/Science_Exploration/Space_Science/Webb/Webb_inspects_the_heart_of_the_Phantom_Galaxy.

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