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Emission from free neutron layer in binary neutron star merger

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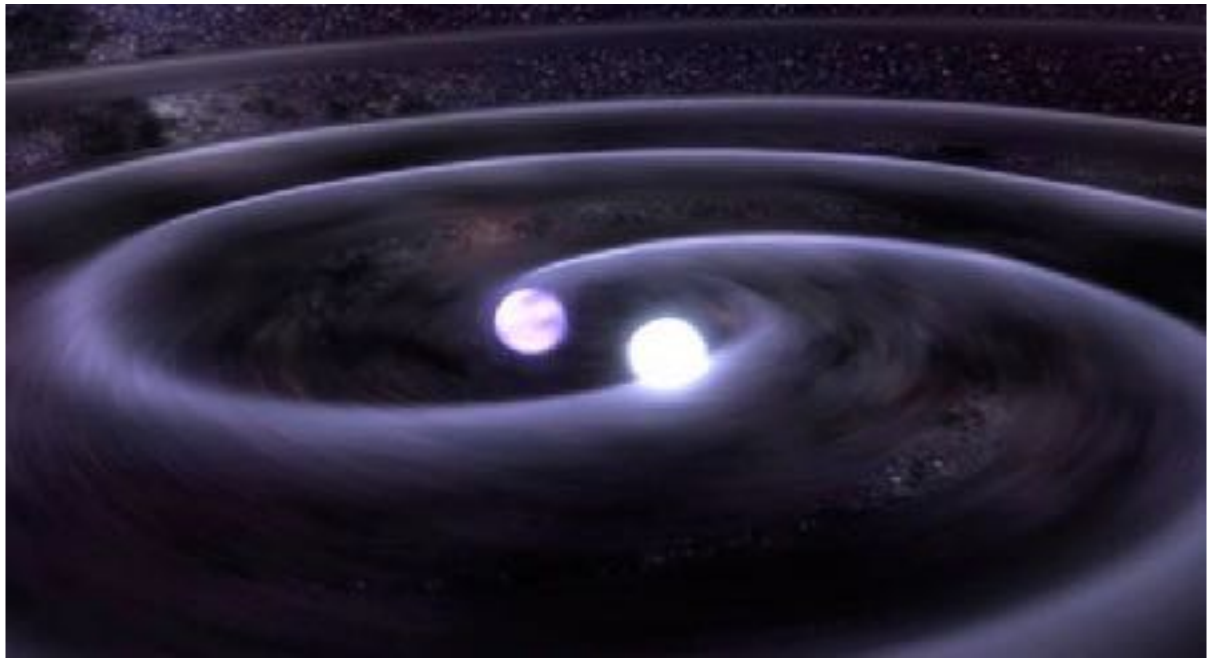
Ishii, TS, Tanaka 2018, arXiv: 1805.04909

Emission from Binary Neutron Star Merger

(Li & Paczynski 1998, B. D. Metzger et al. 2010, ⋯)

GW170817

Electromagnetic emission was detected over wide wavelength range



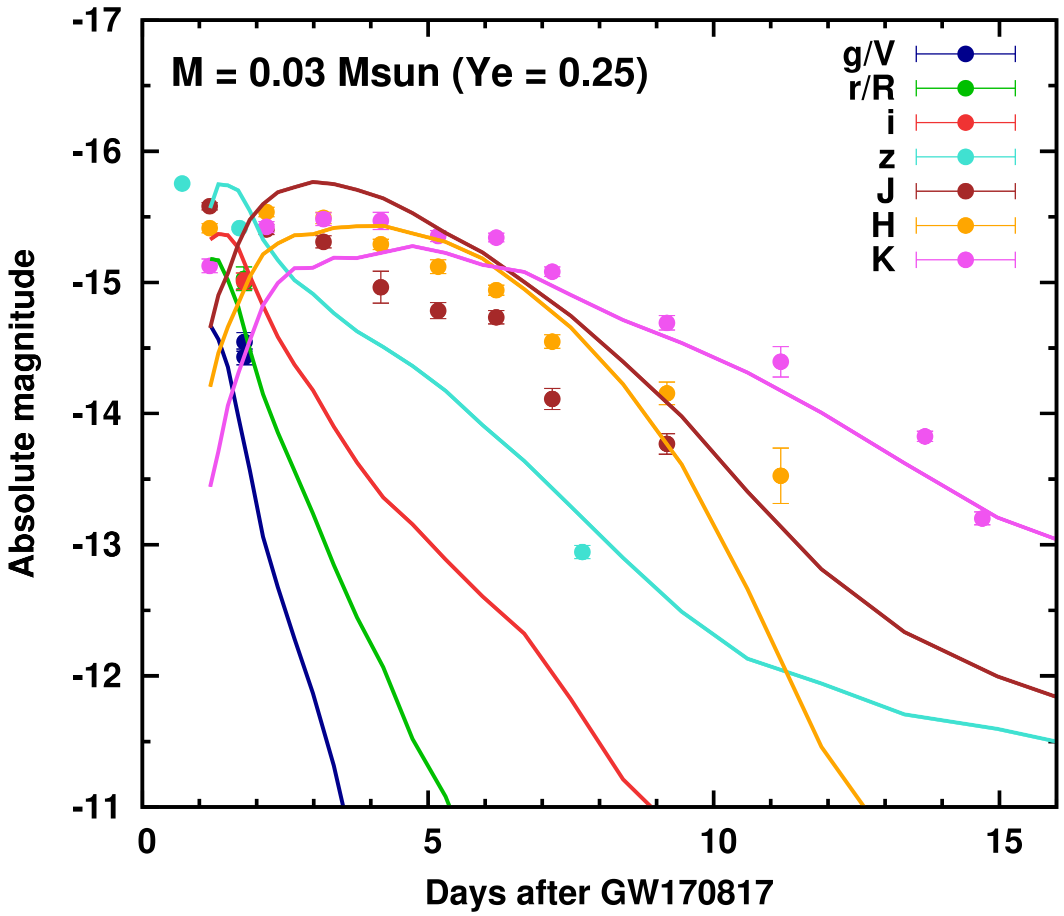
http://aasnova.org/2015/10/28/what-do-you-get-when-two-neutron-stars-merge/

<u>Observed emission could be almost explained by kilonova model</u>

- Neutron star matter is ejected at merging (Neutron-rich ejecta best site for r-process nucleosynthesis)
- Emission from ejecta composed of radioactive elements is detected

Early Emission from Neutron Star Merger

(M. Tanaka et al. 2017, Y. Utsumi et al. 2017, I. Arcavi 2018, ⋯)



- Observations started in ~11 h after merging event
- Observed early emission (~1 day) is more luminous and bluer than model computation

Early emission can provide us with rich information

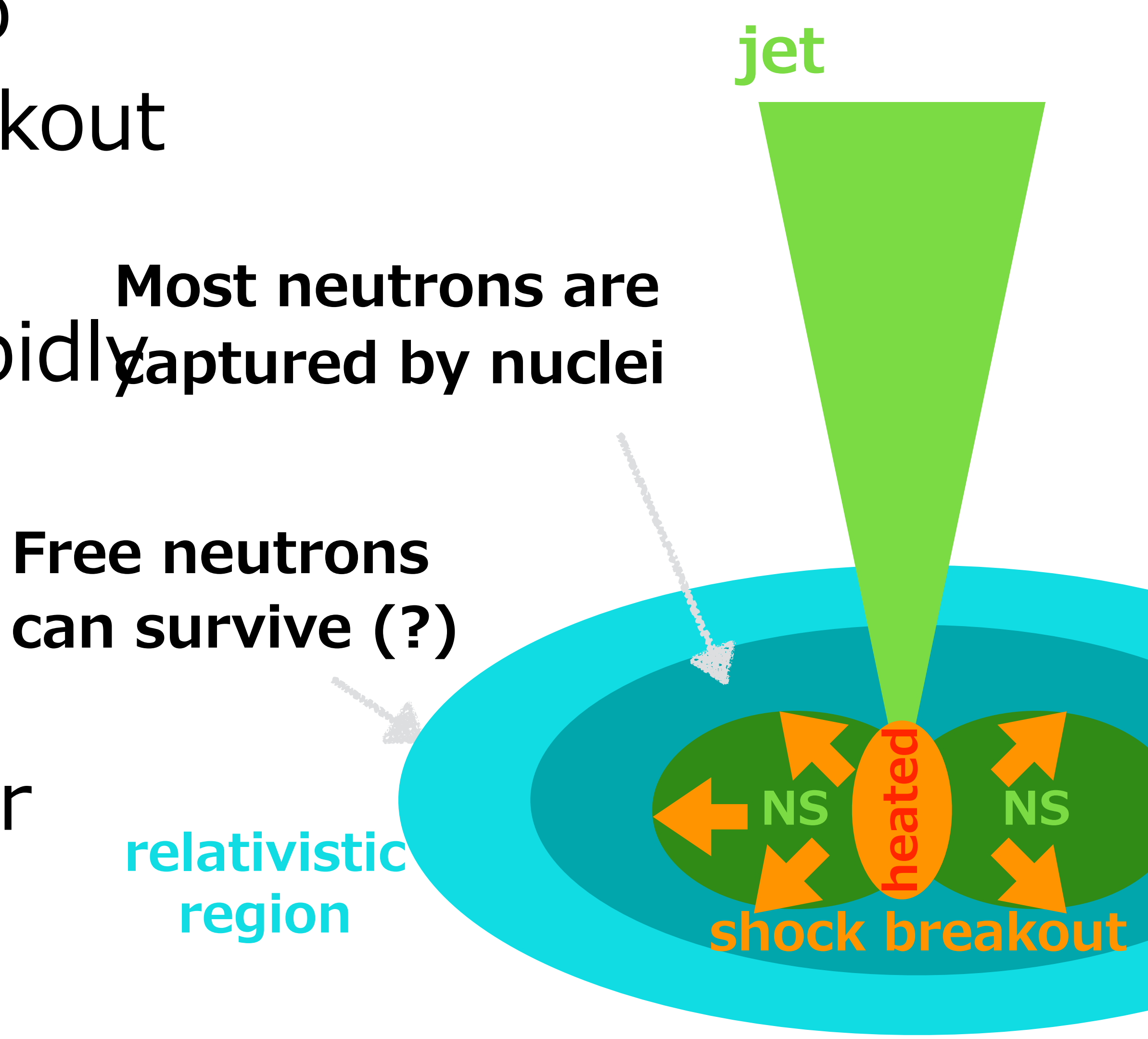
→ Shock produced by the merger may contribute

Free Neutron Precursor

(B. D. Metzger et al. 2015, Metzger 2017)

- Outermost ejecta accelerated to relativistic speeds in shock breakout
  (K. Kyutoku et al. 2014)
- Outermost ejecta expand so rapidly that neutrons avoid capture ($M \sim 10^{-4}\ M_{sun}$)
  (Goriely et al. 2014, Just et al. 2014)
- β-decays of free neutrons power “precursor” to kilonova

peaks at ~ few hours



Most neutrons are captured by nuclei

Free neutrons can survive (?)

Smoothed Particle Hydrodynamics (SPH) simulation (Just et al. 2015)
→Can the similar result be obtained in grid-based simulations?

Objectives

Examining early emission by free-neutron-powered precursor
in shock breakout of binary neutron star merger

Step 1

- Developing relativistic Lagrangian hydrodynamics code and reproducing shock breakout of neutron star merger

Step 2

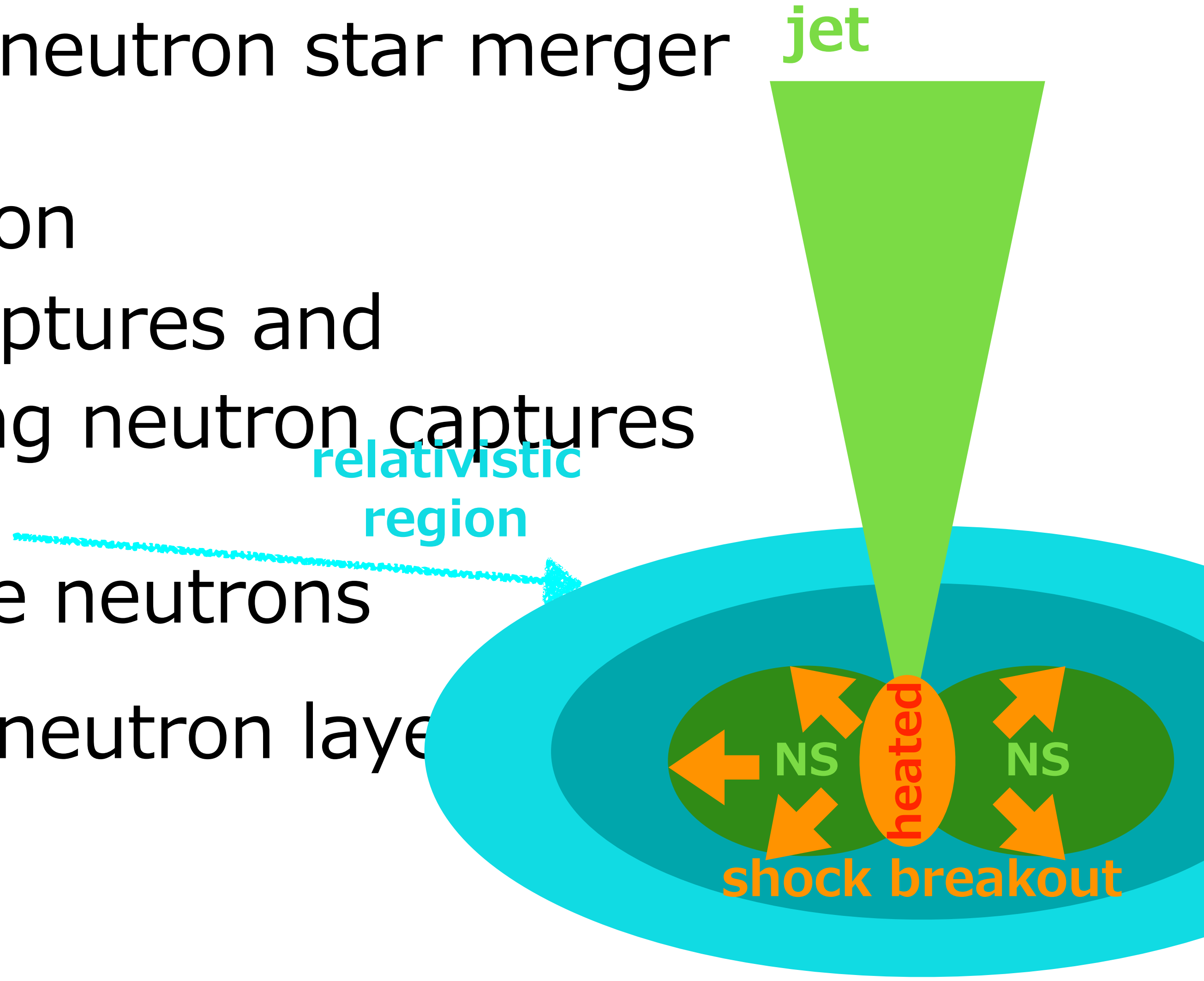
- Estimating surviving free neutron
  mass fraction with e^+ and e^- captures and
  some nuclear reactions including neutron captures

Step 3

- Calculating mass of ejected free neutrons

Step 4

- Calculating emission from free neutron laye



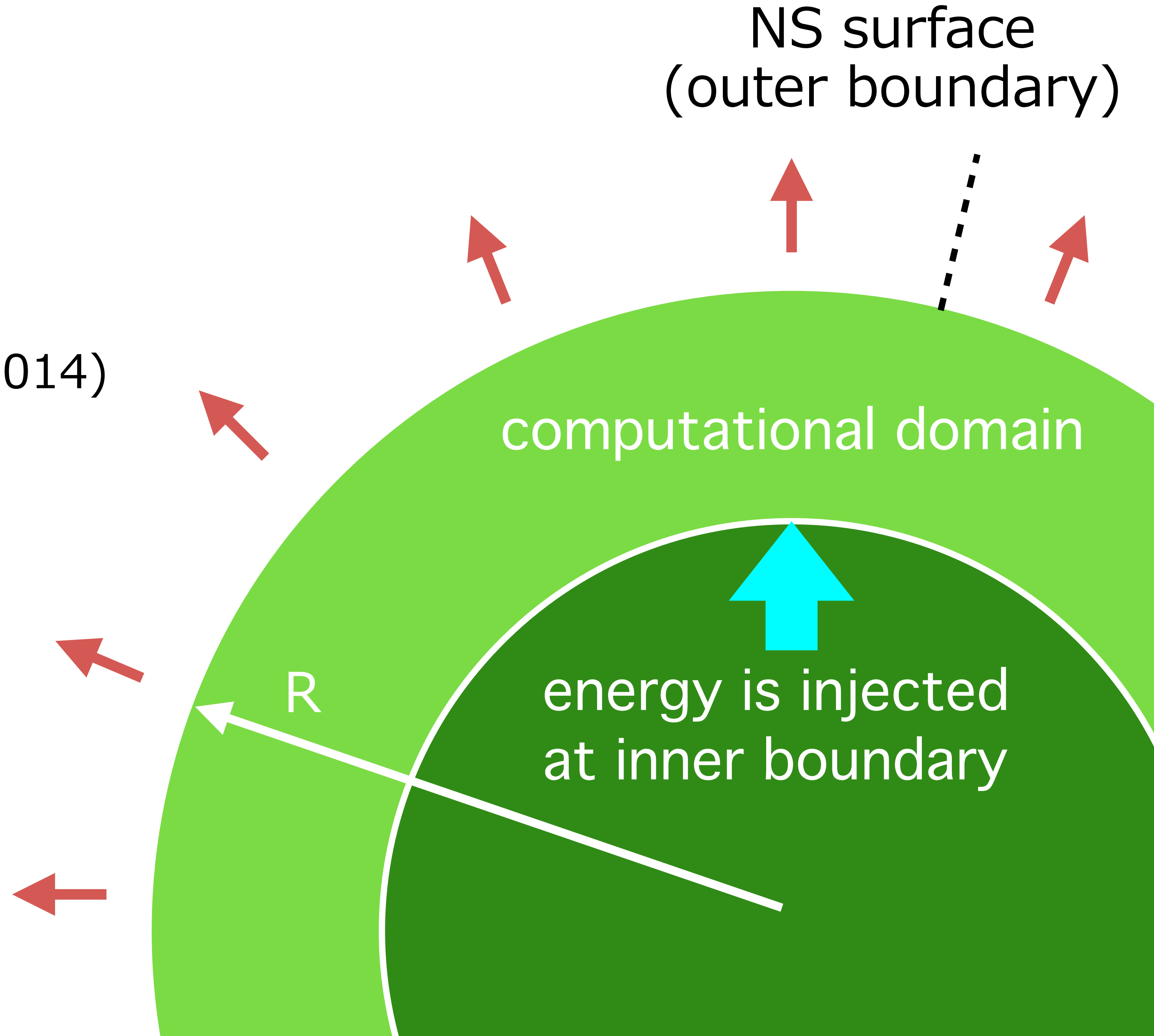
Simulation Condition

- Relativistic Lagrangian hydro simulation
- 1D spherical symmetric coordinate
- 500 computational cells in radial direction
- $E_{\rm final} = 10^{47} - 10^{50}$ erg
- R = 15, 20, 25, 30 km
- $M_{\rm shell} = 10^{-3}\ M_{\rm sun}$
- $\rho \propto (R - r)^3$ (K. Kyutoku et al. 2014)

Shock wave propagates
through merging NS

↓

Shock breakout occurs
when it reaches NS surface



Estimation of free neutron fraction

- Free neutron fraction X_n is set to be 0.9 initially (Y_e=0.1) (beta equilibrium of cold dense matter)
- $e^{\pm}$ is generated by shock heating

$$n + e^{+} \rightarrow p + \bar{\nu_e}$$

$$p + e^{-} \rightarrow n + \nu_e$$

- Time scale of positron and electron capture processes depend on temperature ($\tau_+(T)$, $\tau_-(T)$)

  (L. Kawano 1992, B. D. Metzger et al. 2015)

- Time evolution of X_n when T>10^{10} K is calculated by

$$\frac{dX_\mathrm{n}}{dt} = -\frac{X_\mathrm{n}}{\Gamma\tau_+(T)} + \frac{(1-X_\mathrm{n})}{\Gamma\tau_-(T)}$$

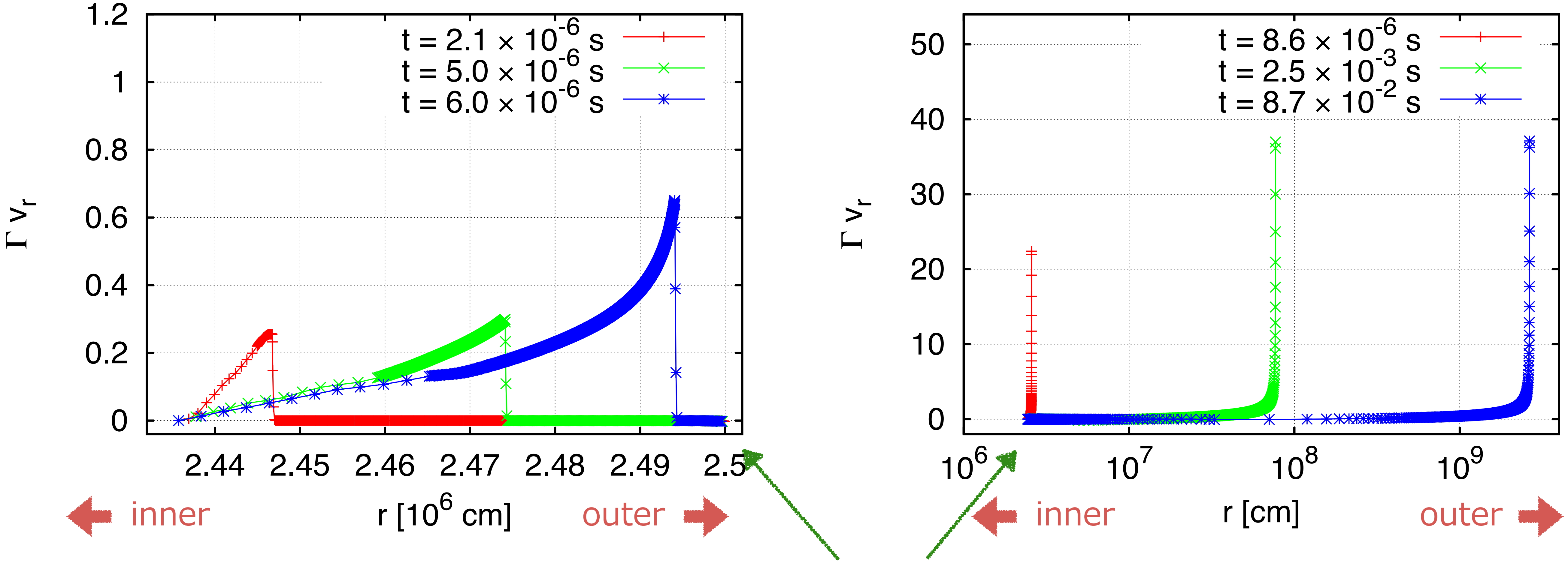
- Nuclear reaction network calculations are performed after temperature decreases down to 10^{10} K (Shigeyama et al. 2010)

Results

Before shock breakout

R = 25 km, E_{final} = 10^{49} erg

After shock breakout



inner

outer

inner

outer

surface

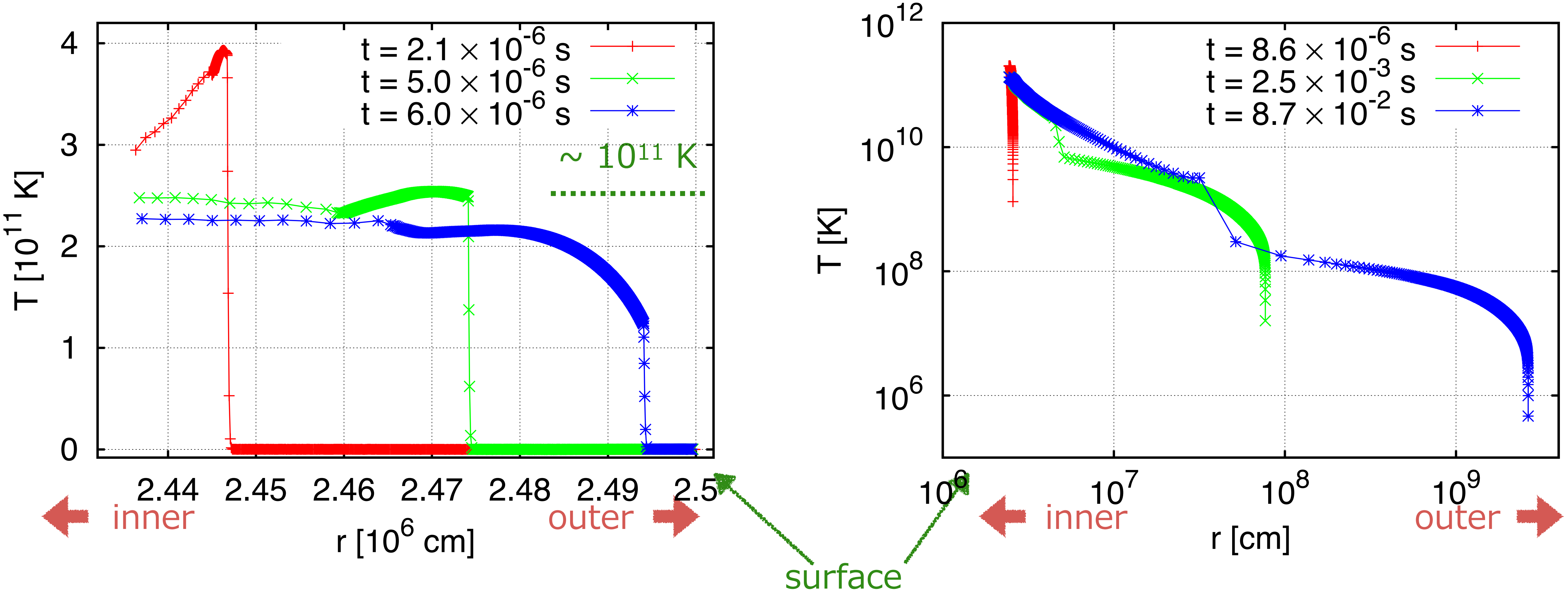
- Accelerated shock wave near the surface can be reproduced
- Ejecta in outermost region have relativistic speeds

Results

Before shock breakout

$R = 25$ km, $E_{\mathrm{final}} = 10^{49}$ erg

After shock breakout



- Radiation pressure is assumed to be dominant ($P = \frac{aT^4}{3}$)
- Temperature decreases rapidly after shock breakout

$$\tau_{+} \simeq 2.1 (T/\mathrm{MeV})^{-5}\ \mathrm{s} \xrightarrow{10^{11}\ \mathrm{K}} \sim 4.41 \times 10^{-5}\ \mathrm{s}$$

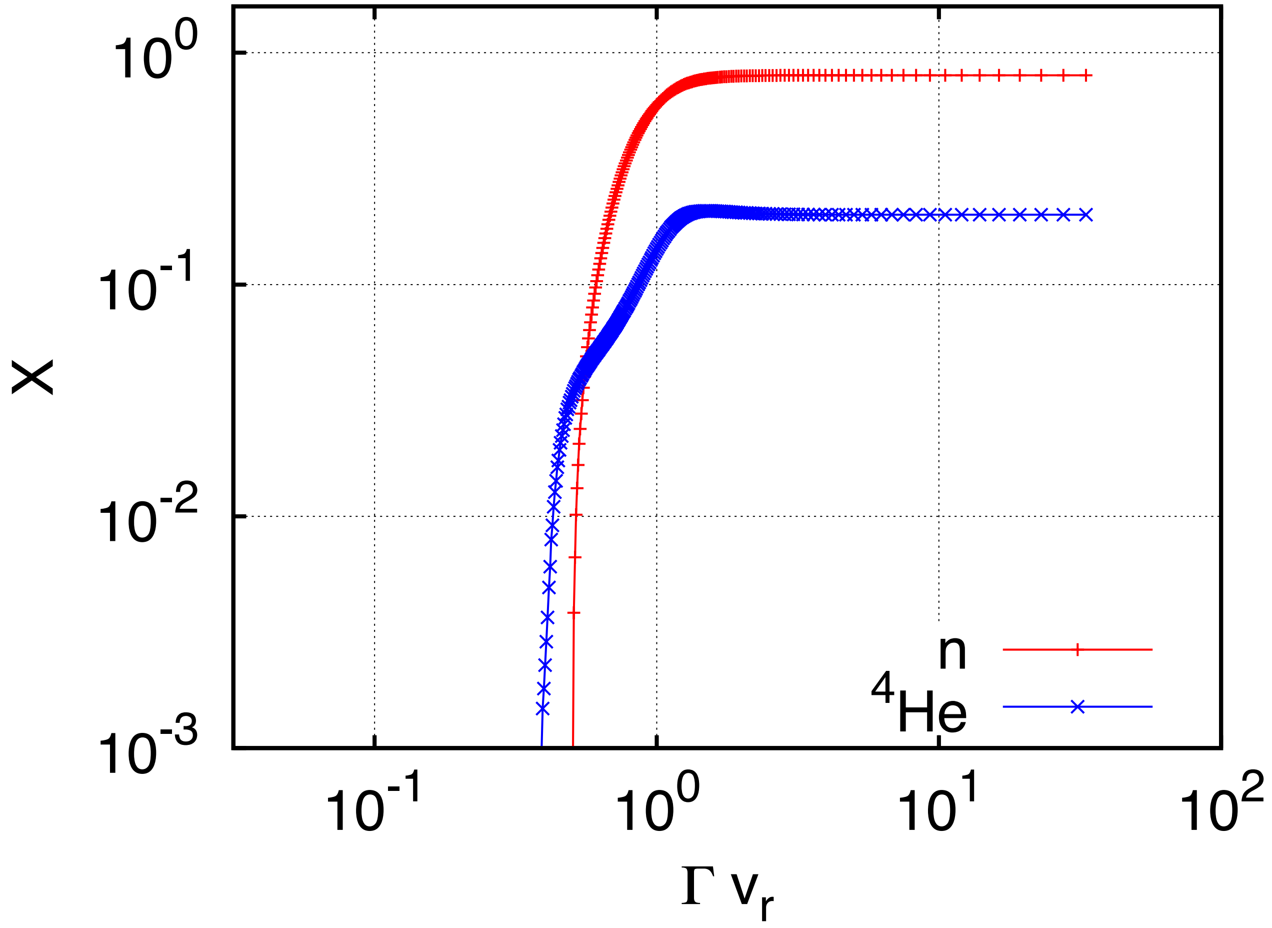
Distribution of Neutrons

R = 25 km

E_{final}= 10^{49} erg

t = 9×10^{-2} s

T < 10^{8} K



Inner region

- Neutrons are captured by nuclei to produce heavy elements

Middle region

- p(n,γ)d reactions consume all neutrons to produce ^{4}He

Outermost region

- Free neutron layer is produced due to low density and temperature

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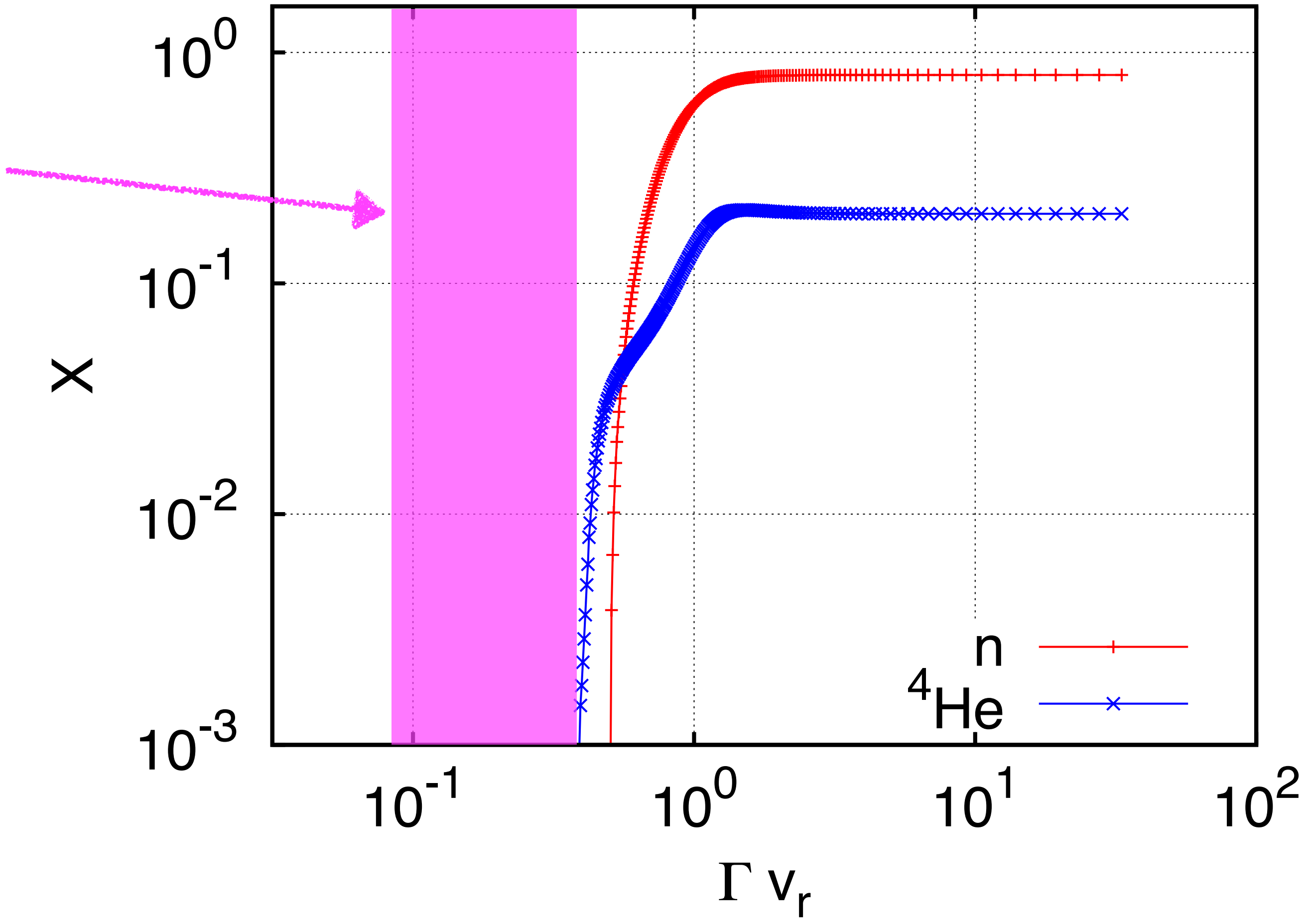
inner
region



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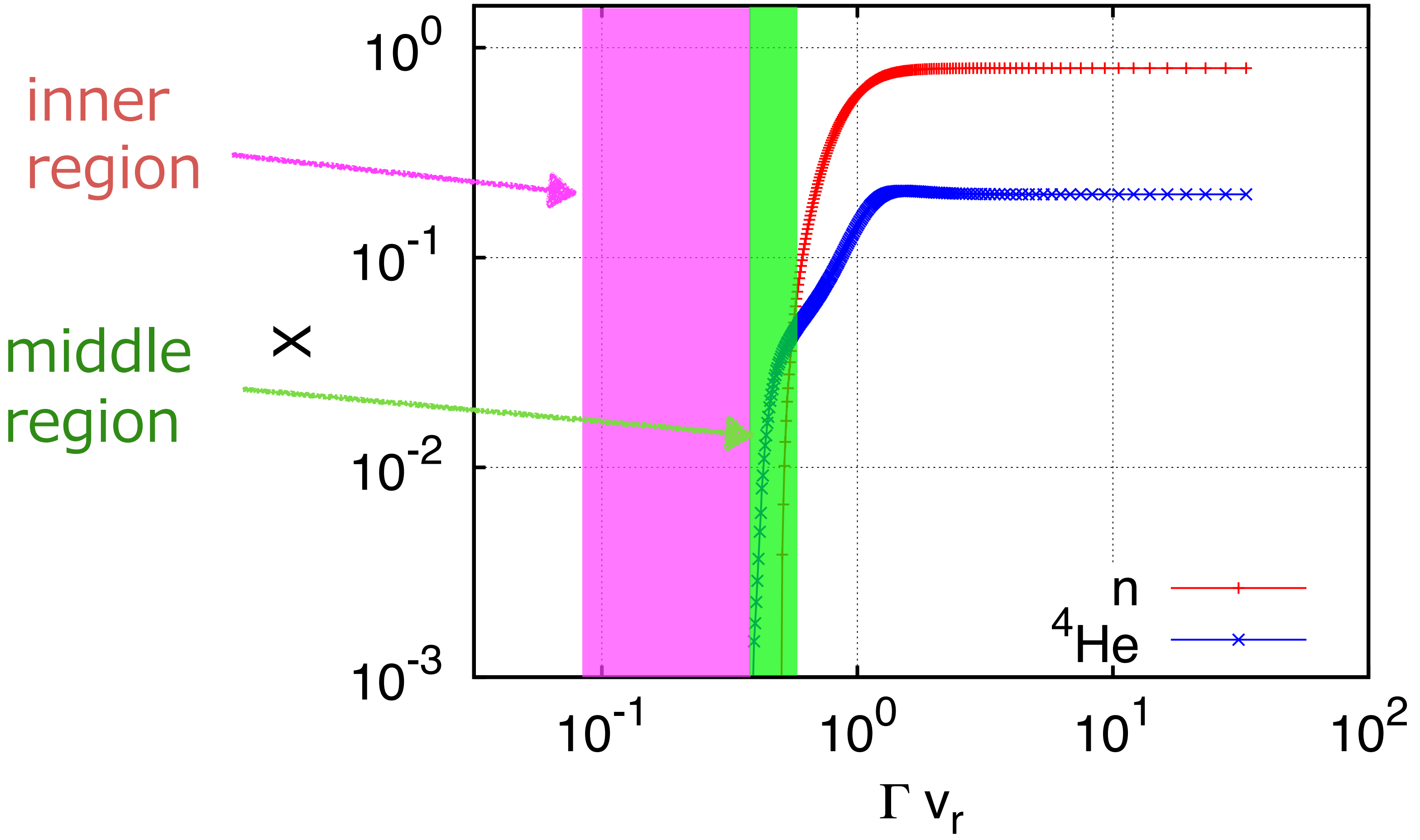

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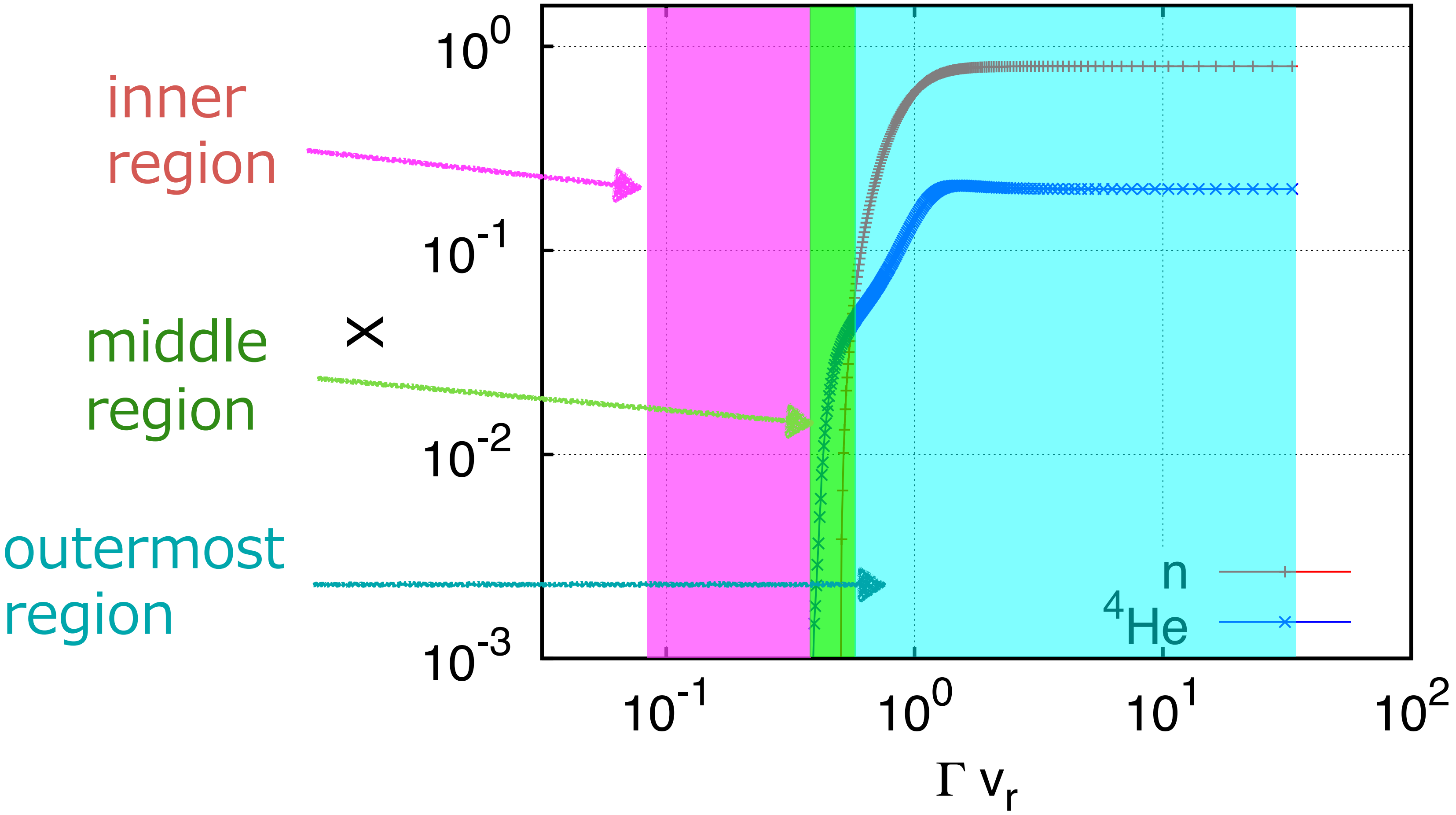

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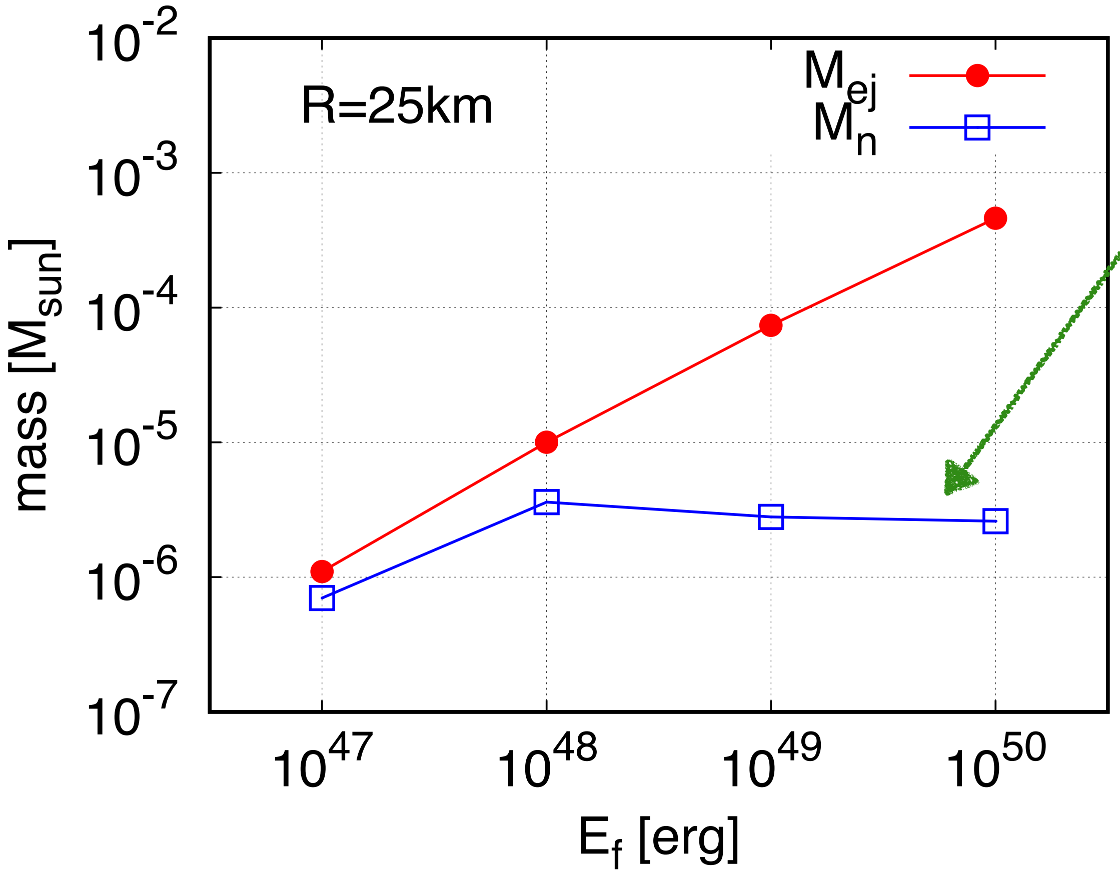
Total Mass of Free Neutrons



$M_n = \Sigma_i (X_{n,i} \times m_i)$

(Total mass of free neutron)

- Models with an energy of 10^{48} erg yields maximum amounts of free neutrons for all *R*
- M_n value is smaller than previous SPH work by ~2 orders of magnitude (~10^{-4} M_{sun})

R [km] ($E_f = 10^{48}$ erg)	15	20	25	30
M_n $[M_{sun}]$	9.2×10^{-7}	2.1×10^{-6}	3.6×10^{-6}	5.2×10^{-6}

Emission from free neutron layer

M_{ej} = 10^{-5} M_{sun}, E_{final} = 10^{48} erg, ejecta velocity ~c/3, opacity ~0.4 cm^2 g^{-1}

- Photon diffusion velocity becomes comparable to expansion velocity (c/3) at ~1,500 s
- Energy density at neutron decay time (~800 s) is

$$\epsilon_0 \sim 10^6 \ \mathrm{erg\ cm^{-3}} \ (M_\mathrm{n}/3.6 \times 10^{-6} \ M_\mathrm{sun})$$

- Subsequent adiabatic expansion up to 1,500 s, luminosity L is estimated by

$$L \sim 7.6 \times 10^{41} \ \mathrm{erg\ s^{-1}} \ \left(\frac{t}{1{,}500 \ \mathrm{s}}\right)^{-2} \left(\frac{M_\mathrm{n}}{3.6{\times}10^{-6} \ M_\mathrm{sun}}\right)$$

(Ultraviolet, timescale of ~30 min)

- Absolute magnitudes ~ −13.55 and −13.25 mag (AB magnitude) at 2000 and 2600 Å (UVW2 and UVW1 of Swift/UVOT), respectively

Summary

- Shock breakout in neutron star merger was investigated by relativistic Lagrangian hydrodynamics code
- We obtain the maximum amount of free neutrons when $E_f \sim 10^{48}$ erg
- Total mass of neutron surviving region is $\sim 10^{-6}\ M_{sun}$ (two orders smaller than previous SPH work)
  →Free neutrons might be included in different ejecta components
- Luminosity of free neutron emission reaches $\sim 7\times10^{41}$ erg s^{-1} in optical band at ~ 30 min after merger event

Future work

- Monte Carlo Radiative transfer computation (A. Ishii et al. 2017) with thermal photons from free neutron decays

Thank you for your attention!

Reaction timescales

timescale for positron capture

$$\tau_{+} \simeq 2.1 \left(\frac{kT}{\mathrm{MeV}}\right)^{-5} \ \mathrm{s}$$

(B. D. Metzger et al. 2015)

timescale for electron capture

$$\tau_{-} \simeq \frac{\tau e^{qz}}{\left(\frac{5.252}{z} - \frac{16.229}{z^2} + \frac{18.059}{z^3} + \frac{34.181}{z^4} + \frac{27.617}{z^5}\right)}$$

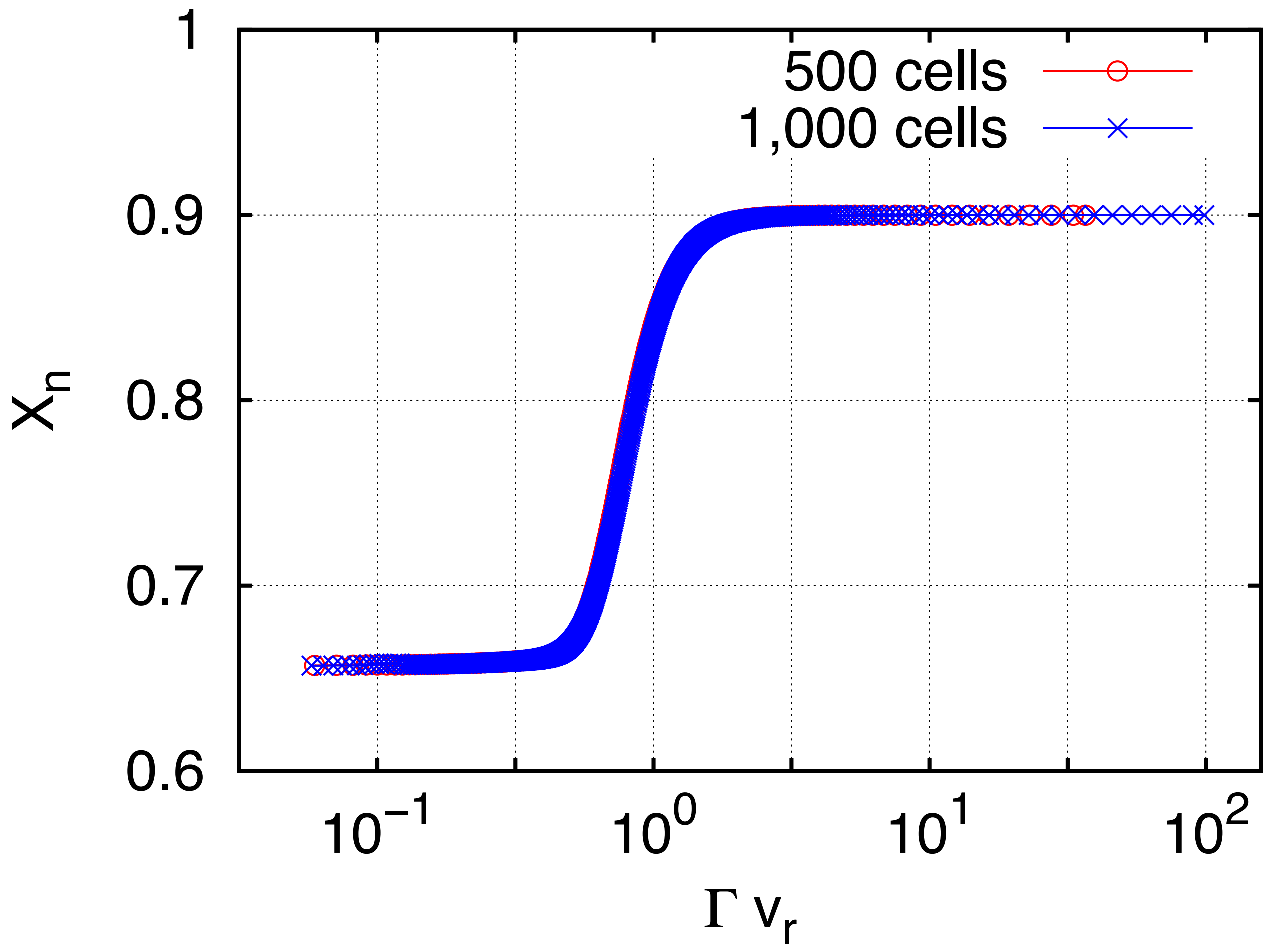
τ: neutron lifetime

q: $(m_n - m_p)/m_e$

z: $m_e c^2/kT$

(L. Kawano 1992)

Grid Convergence Test



- Computation was performed with 1,000 cells
- Distribution extends to larger Γv_r region than 500 cells
- M_n values are almost equivalent
  ($1.4 \times 10^{-6}\ M_{sun}$ and $1.3 \times 10^{-6}\ M_{sun}$)
  →Computation with 500 cells are converged

Endothermic reaction

- Disintegration for the heavy element nuclei is endothermic reaction
- The energy density: $a_r T^4/\rho \sim 3 \times 10^{19}$ erg g^{-1} ($T \sim 2.5 \times 10^{11}$ K)
- The energy of the endothermic reaction for Fe: ~9 MeV per nucleon
  →Energy density for the reaction: 9 MeV/$m_u \sim 8.6 \times 10^{18}$ erg g^{-1}
- Thus the energy is brought out by the endothermic reaction by a few tens of a percent
- The expelled energy might be decreased with the realistic composition of the merging neutron stars